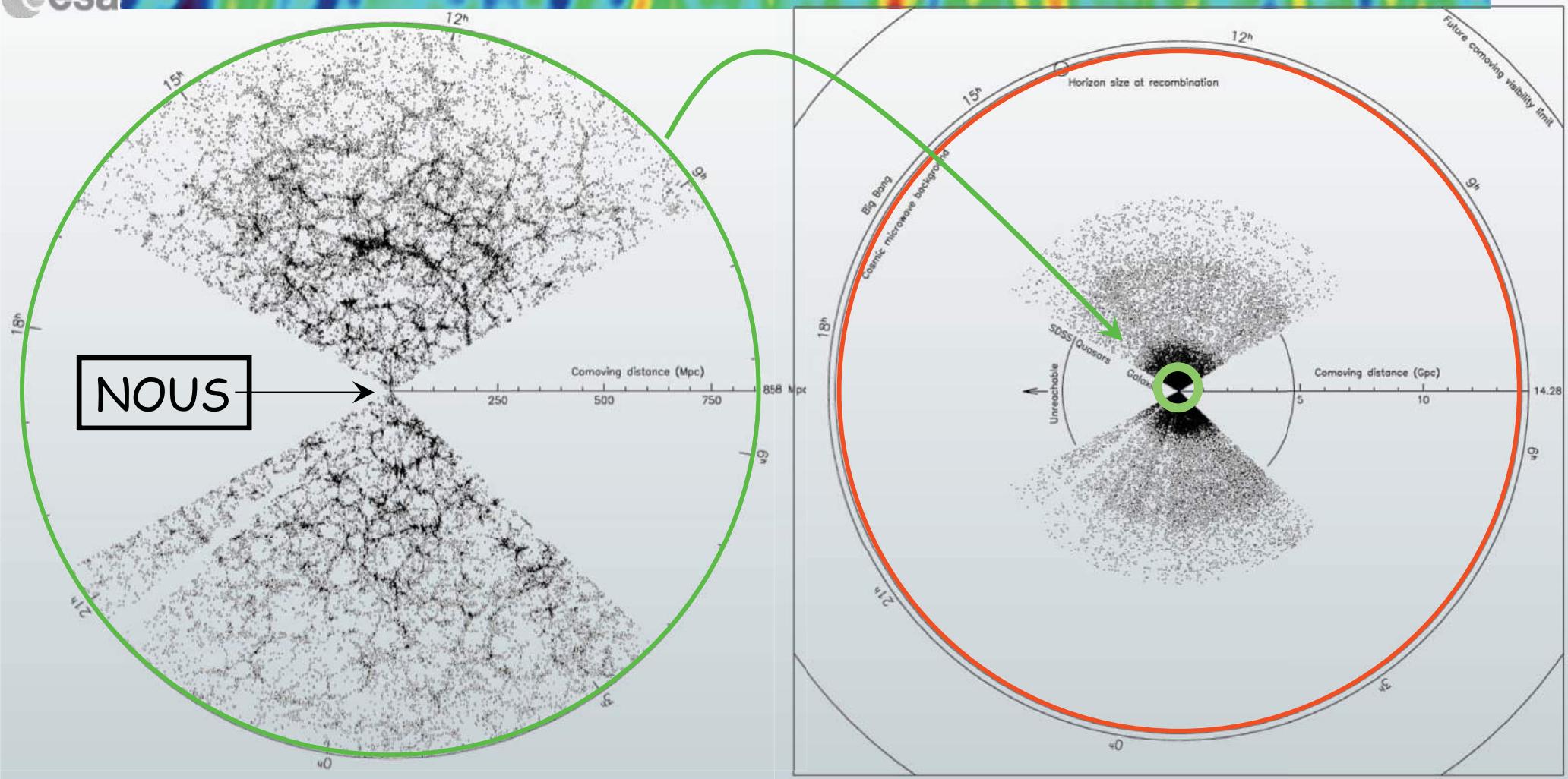


# Archeologie Cosmologique Et lumière primordiale

François R. Bouchet  
Institut d'Astrophysique de Paris

# Voir loin, c'est voir le passé lointain !



Chaque point est une galaxie comme la Notre. La plus proche, M31, est à  $\sim 2.5$  Mpc. Il faut 2,7 milliards d'années à la lumière d'une galaxie sur le cercle vert pour qu'elle nous parvienne.

La lumière millimétrique a mis environ 13 milliards d'années pour nous parvenir (cercle rouge). C'est la trace (comme fossilisée) de la fournaise primordiale, 400 000 ans après le Bang, quand l'Univers est devenu Transparent.

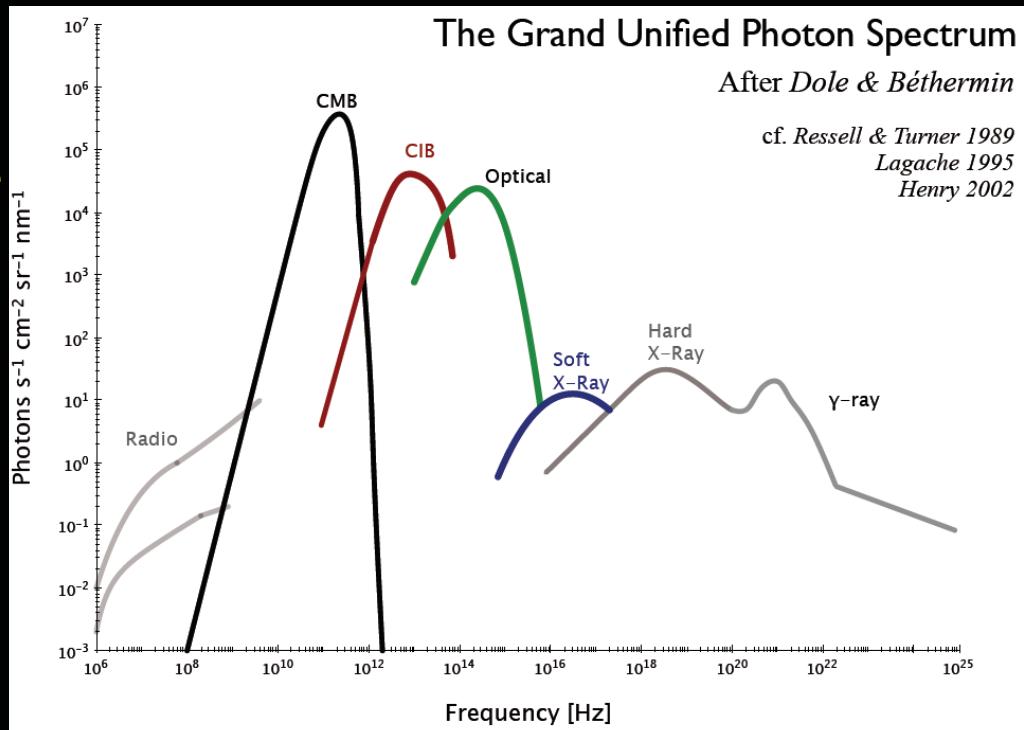


Les photons du rayonnement fossile sont les plus nombreux.

*Il est donc facile de voir le big bang.*

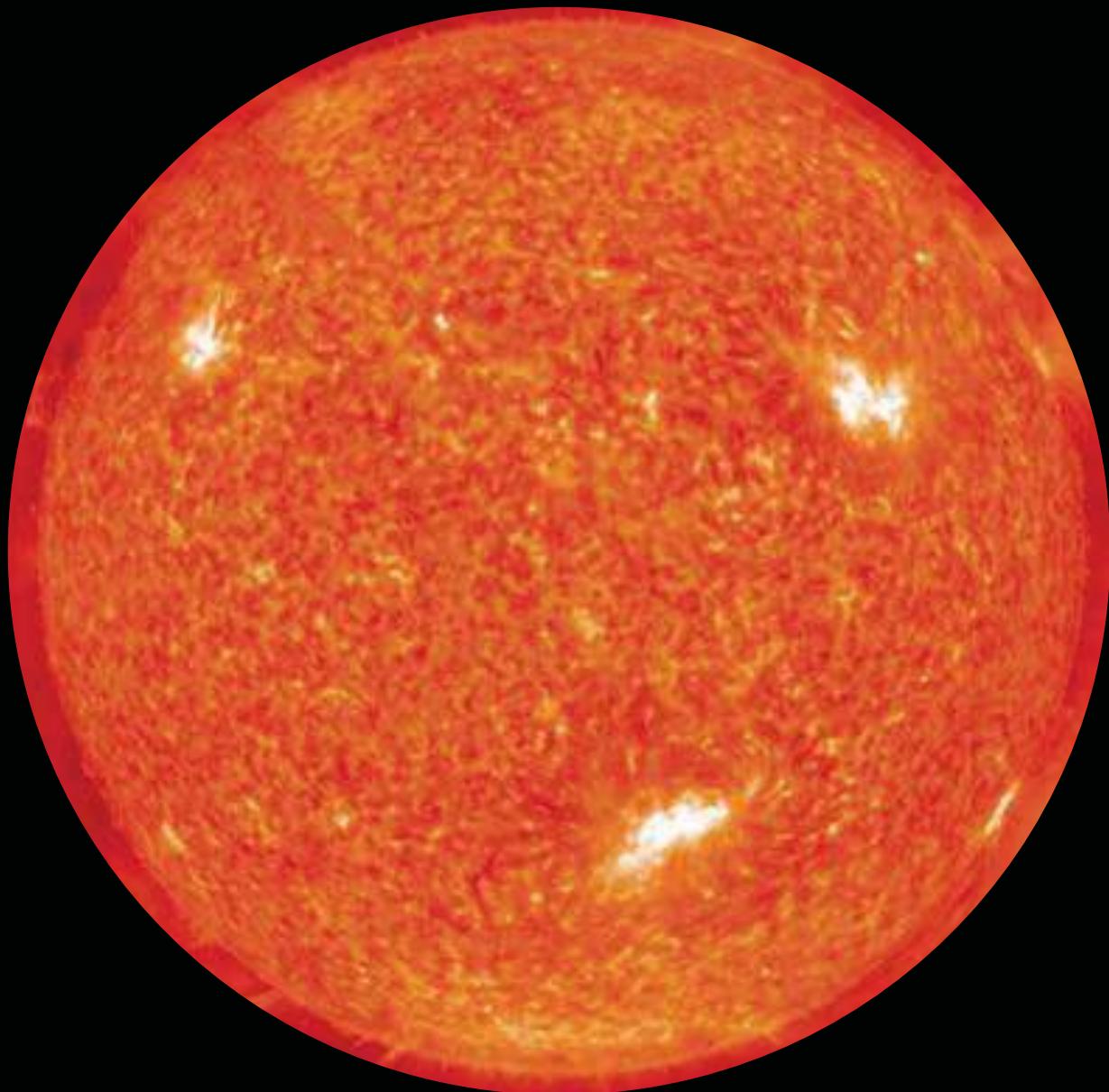
Pas besoin d'un telescope géant, il suffit de regarder là où il faut !

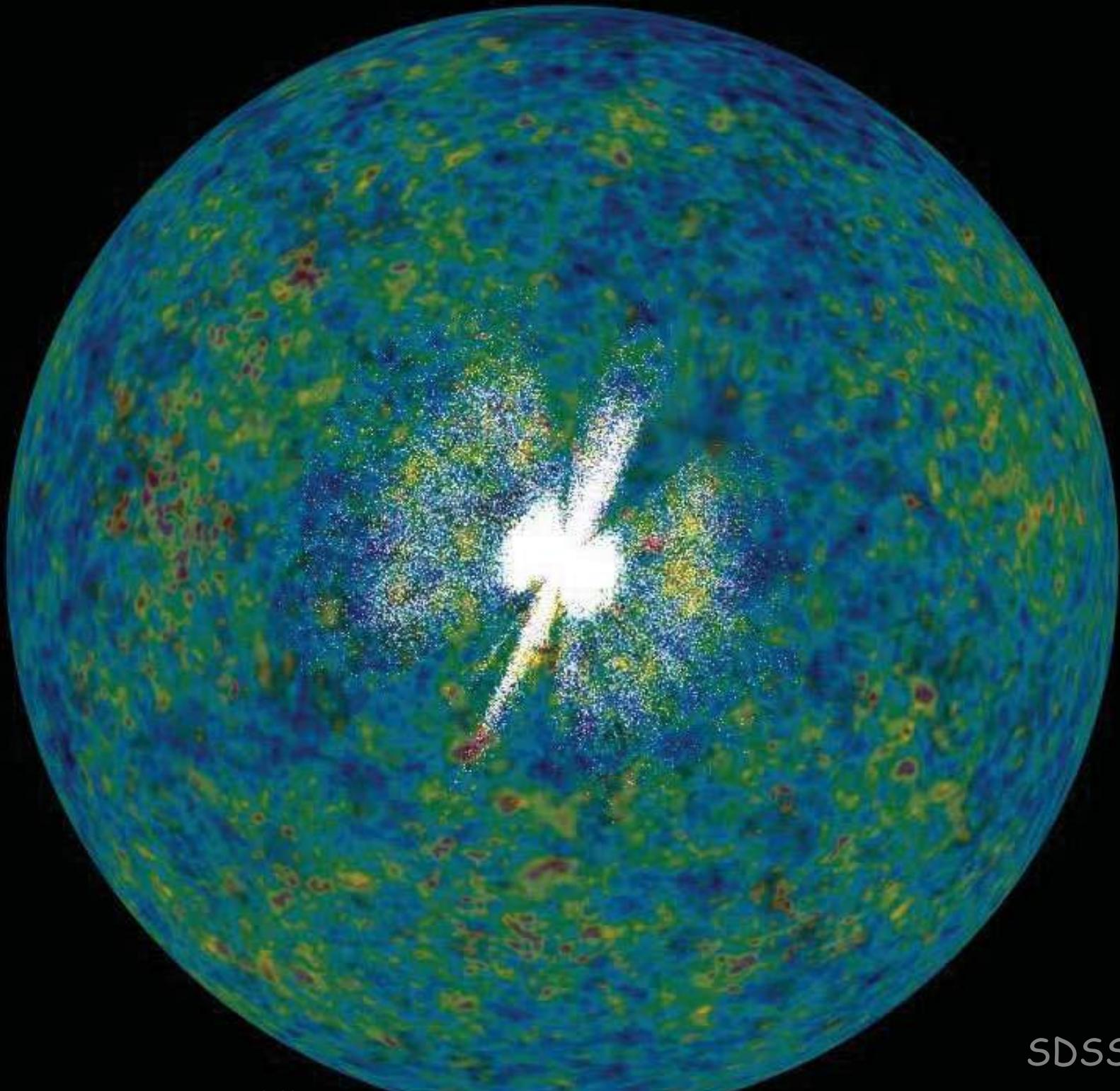
Nous sommes baignés par la lumière du big bang. À cause de l'expansion, elle est décalée vers le rouge dans les micro-ondes et les ondes radios.  
On peut voir ce rayonnement sur un poste de télévision mal réglé !!!



# Une autre surface des dernières diffusions...

- Le soleil est aussi un plasma.
- On ne voit que sa surface, l'intérieur est opaque.
- Les photons qui nous parviennent du soleil n'interagissent plus avec le plasma, il diffusent une dernière fois, formant la surface du soleil avant de nous parvenir.
- En étudiant la surface du soleil, on déduit son fonctionnement à l'intérieur, bien qu'on ne puisse l'observer directement



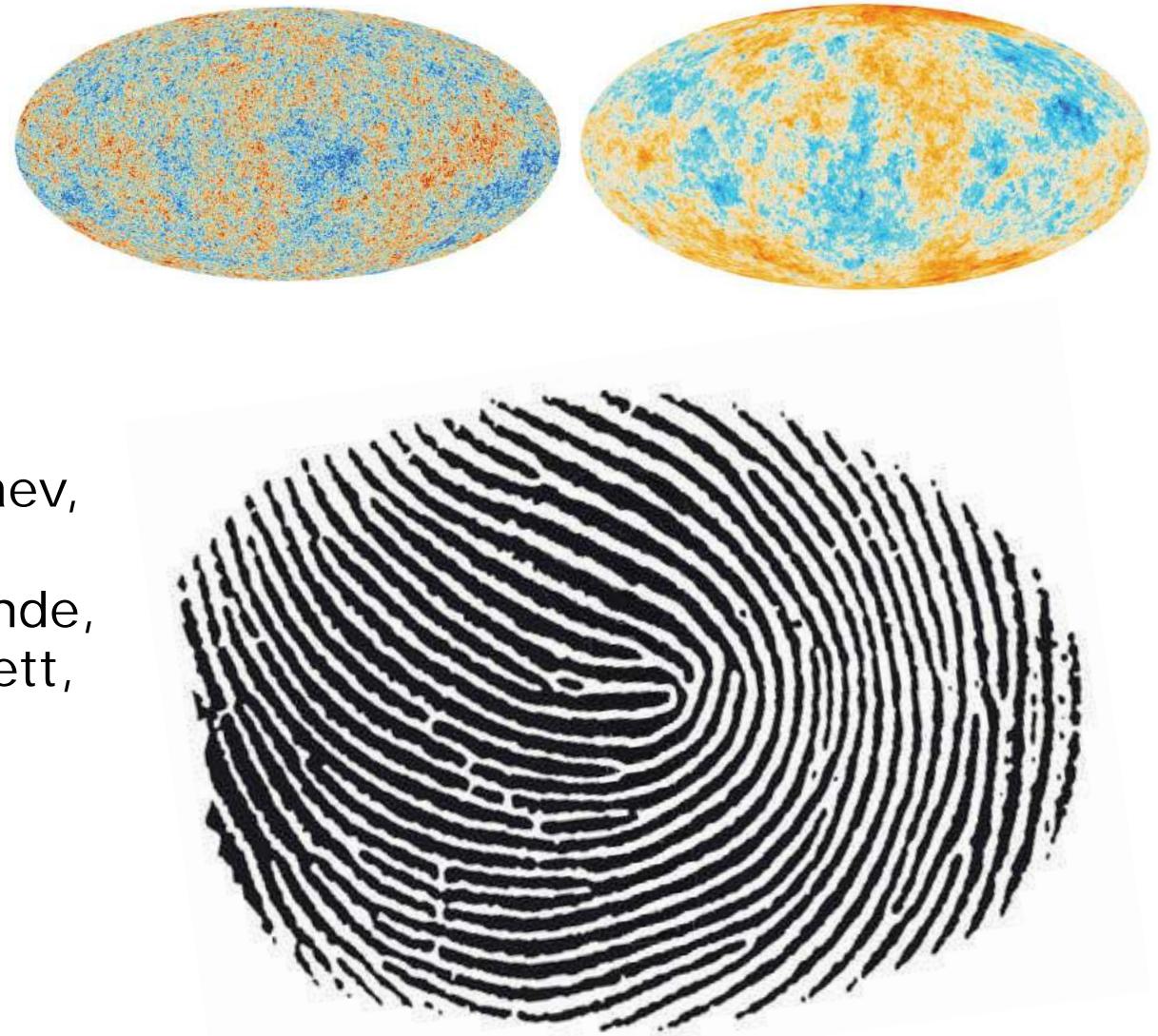


SDSS & WMAP

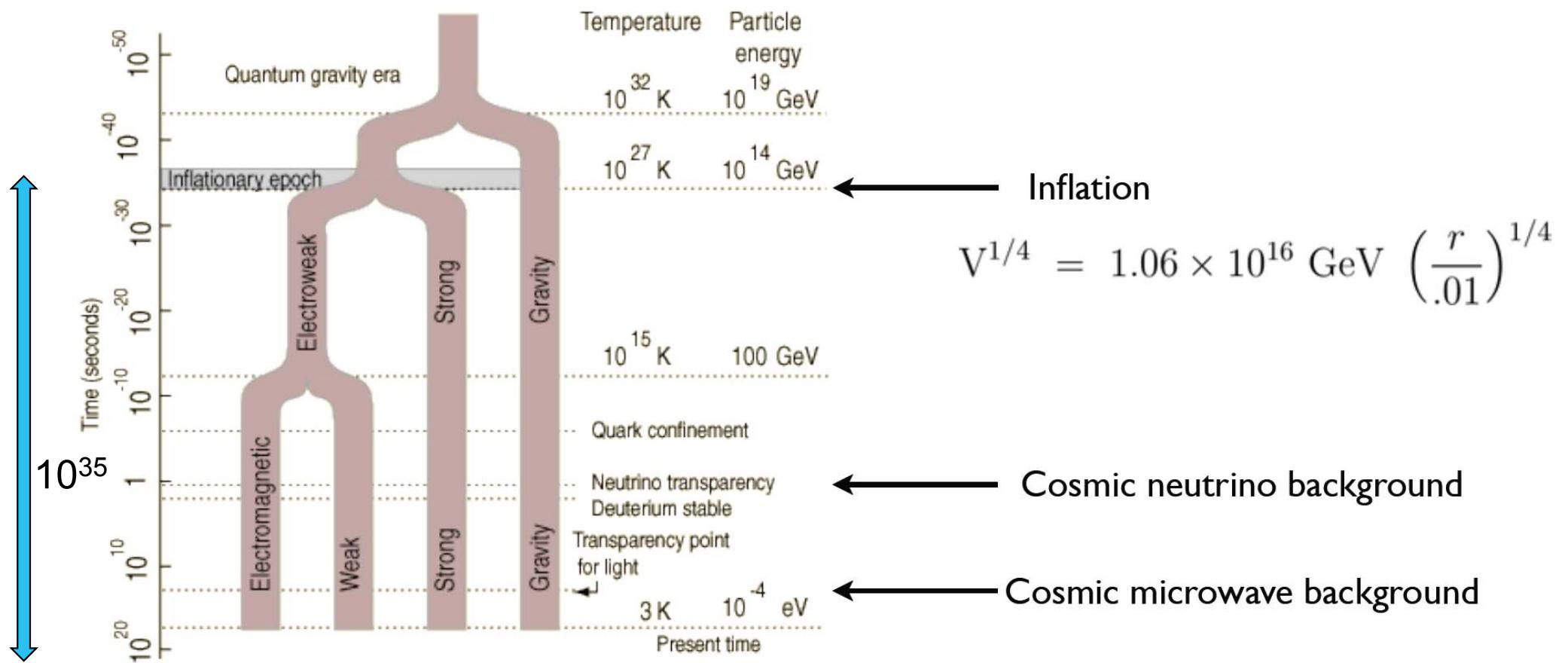
# Les théoriciens ont pré-calculé les empreintes attendues dans le cas de divers scenarii



Peebles, Yu, Sachs & Wolf, Sunyaev, Zeldovich, Silk, Vittorio, Wilson, Mukhanov, Chibisov, Bardeen, Linde, Bond, Efstathiou, Bouchet, Bennett, Gott, Kaiser, Stebbins, Allen, Shellard, Seljak, Zaldarriaga, Kamionkowski, Hu, ...



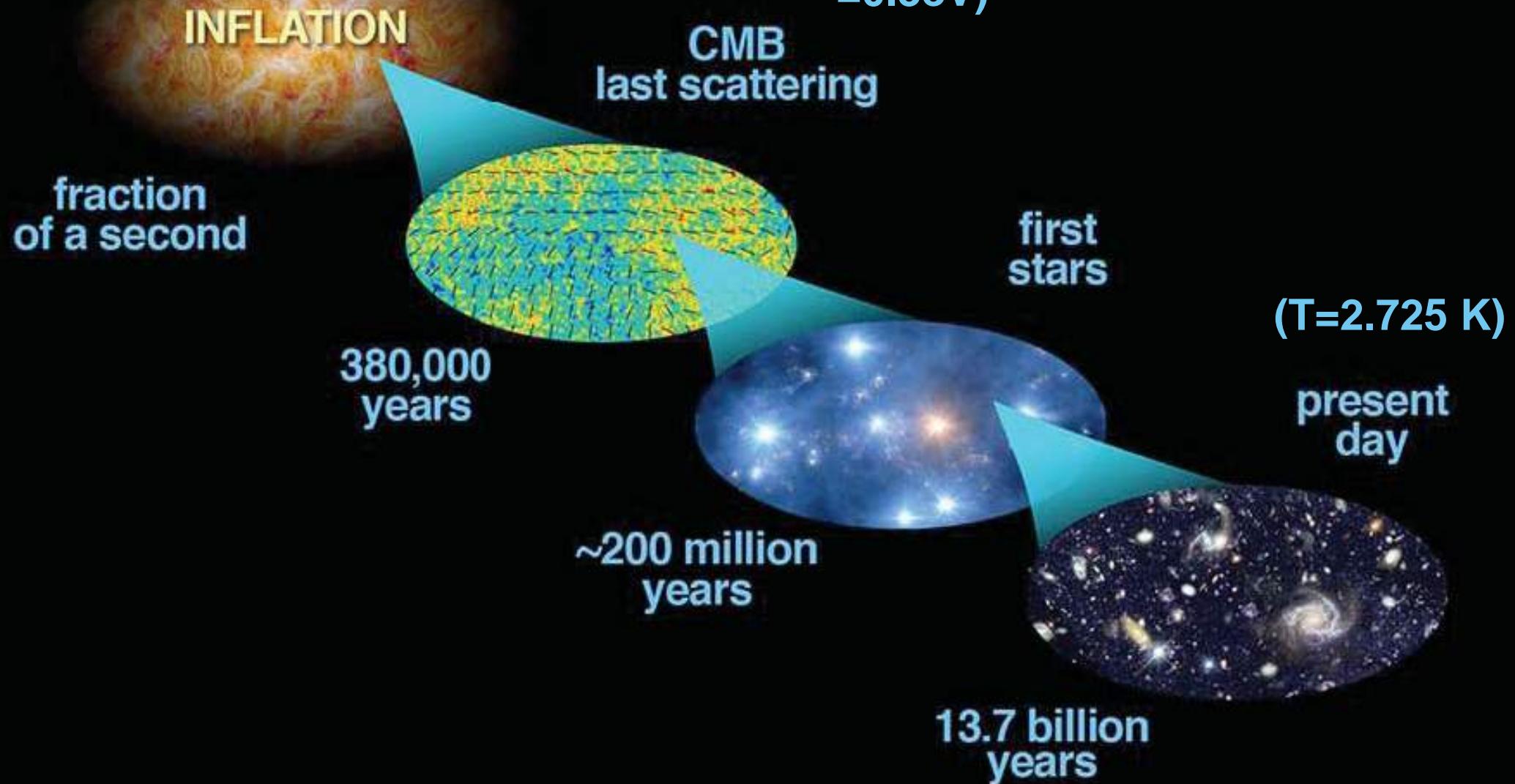
# L'Univers primordial, laboratoire ultime de physique fondamentale



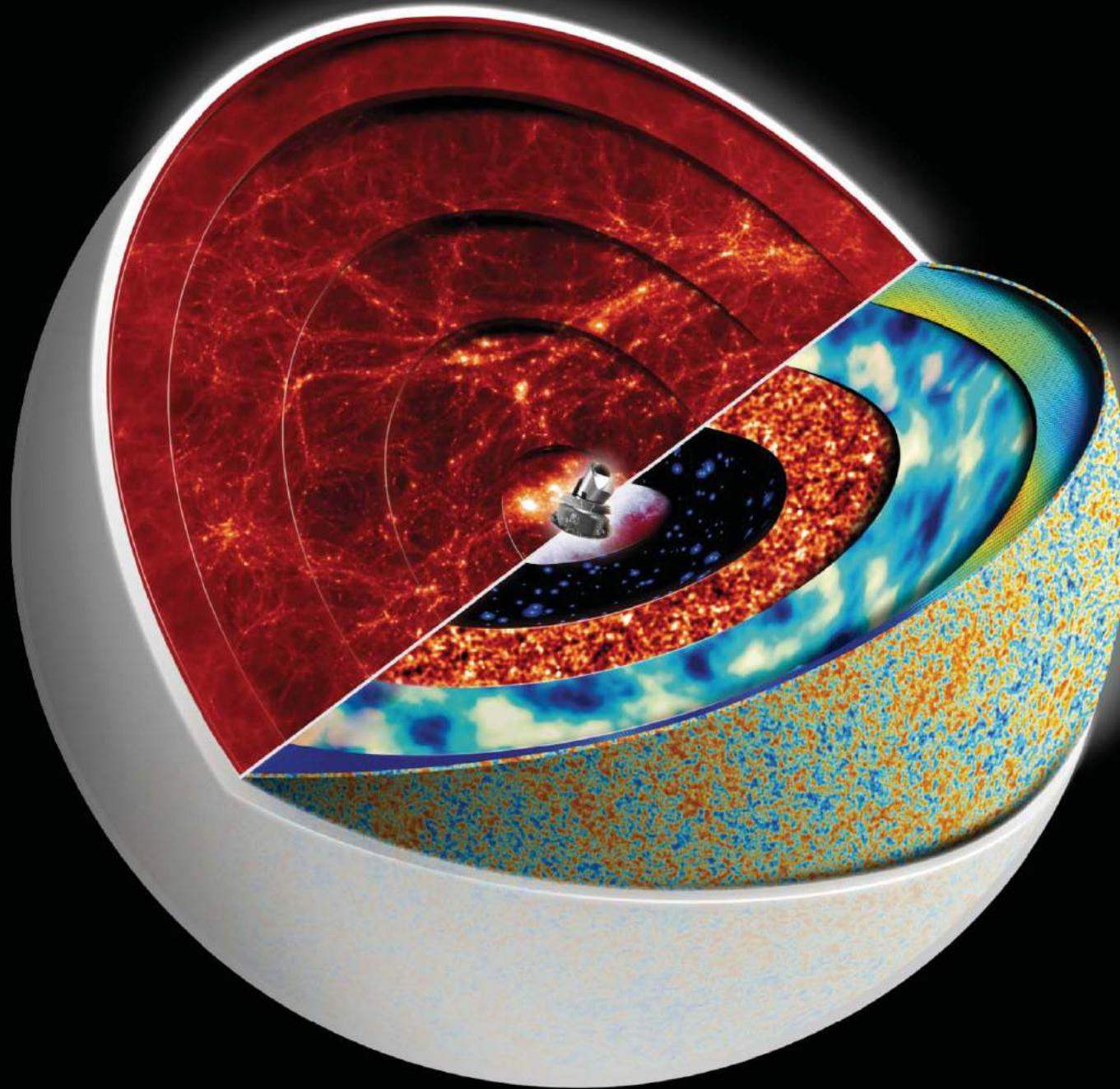
# *Planck* era

( $T=10^{25}$ eV ?)

( $T=3000$  K  
 $=0.3$ eV)



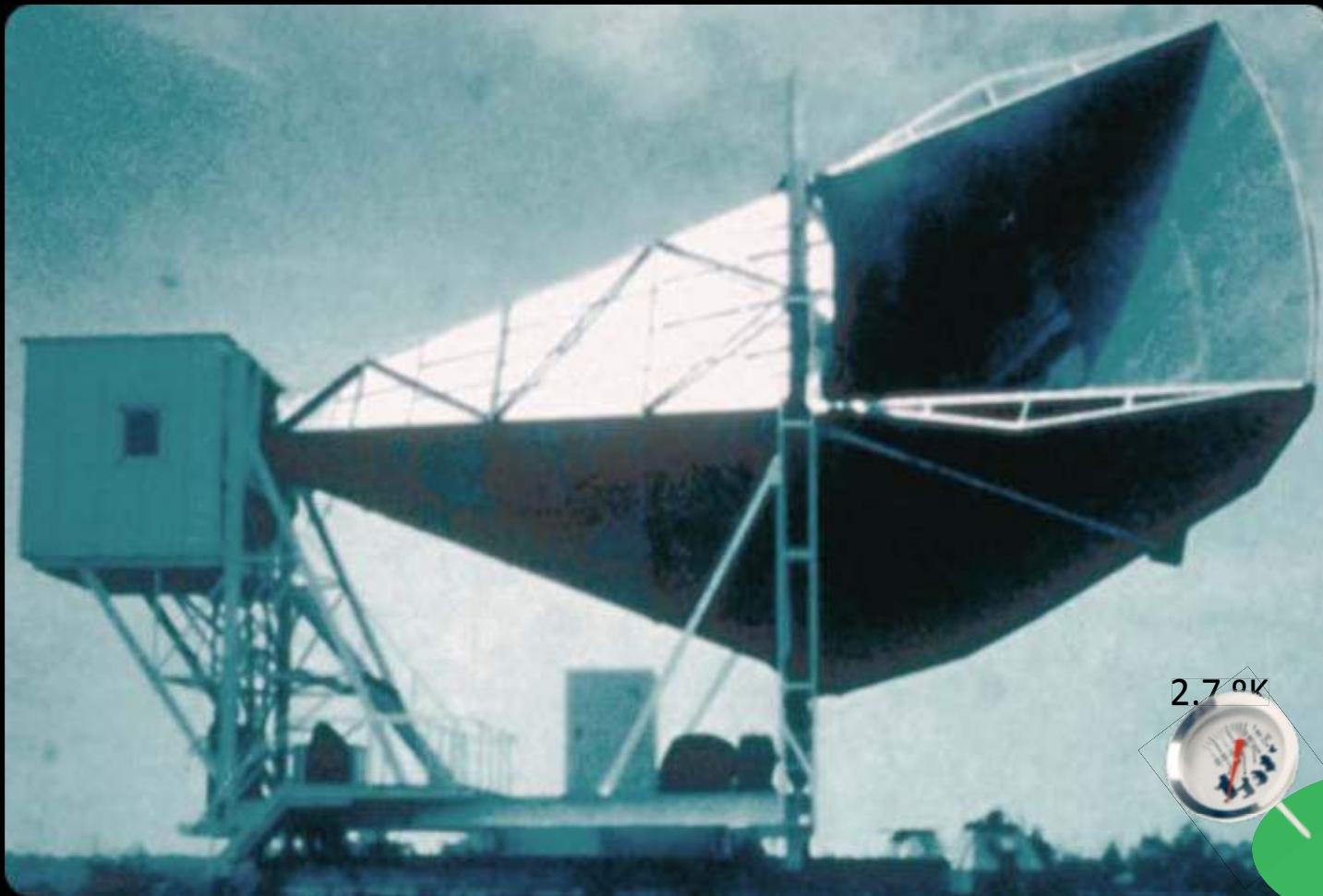
# De la mousse quantique à l'éponge cosmique



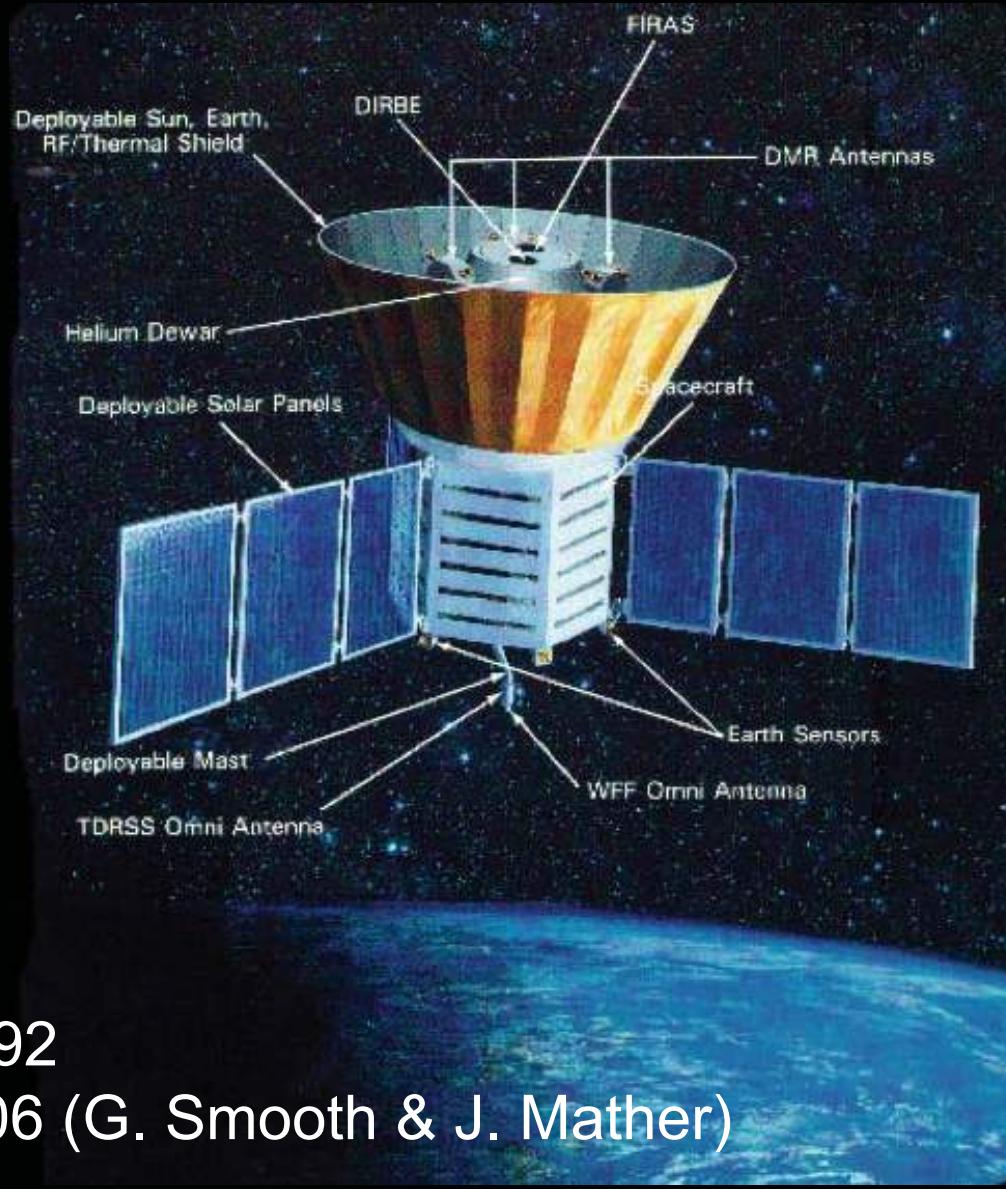


# OBSERVER LE RAYONNEMENT FOSSILE

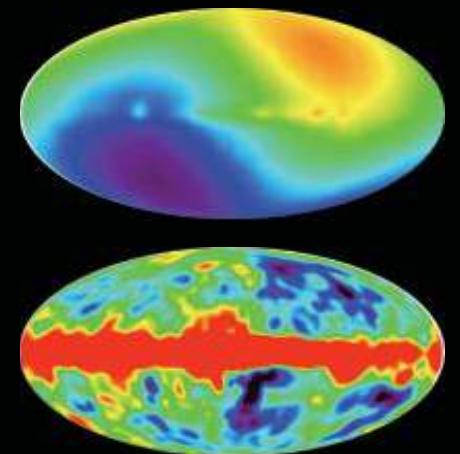
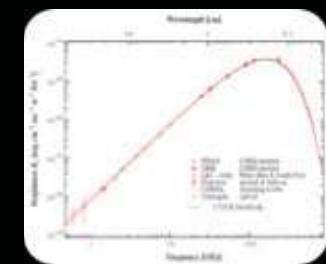
Quelques temps forts

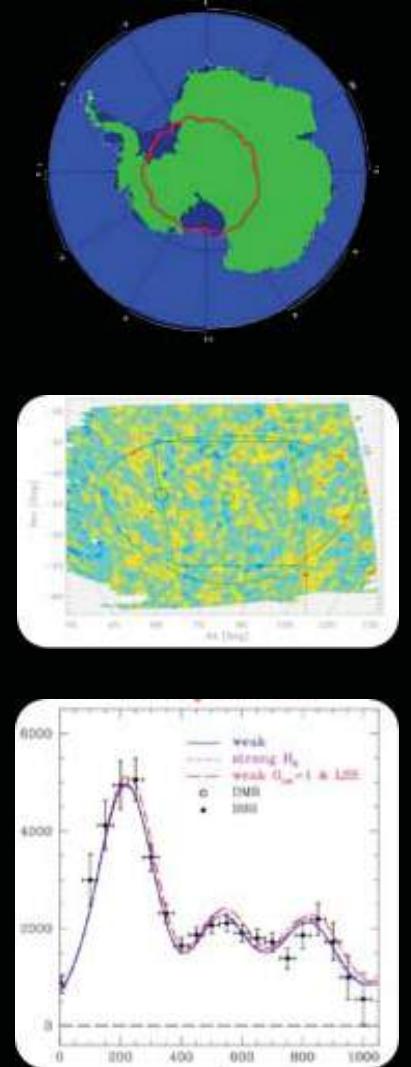


First observation at Bell Labs  
A. Penzias & A. Wilson in 1965  
Nobel 1978

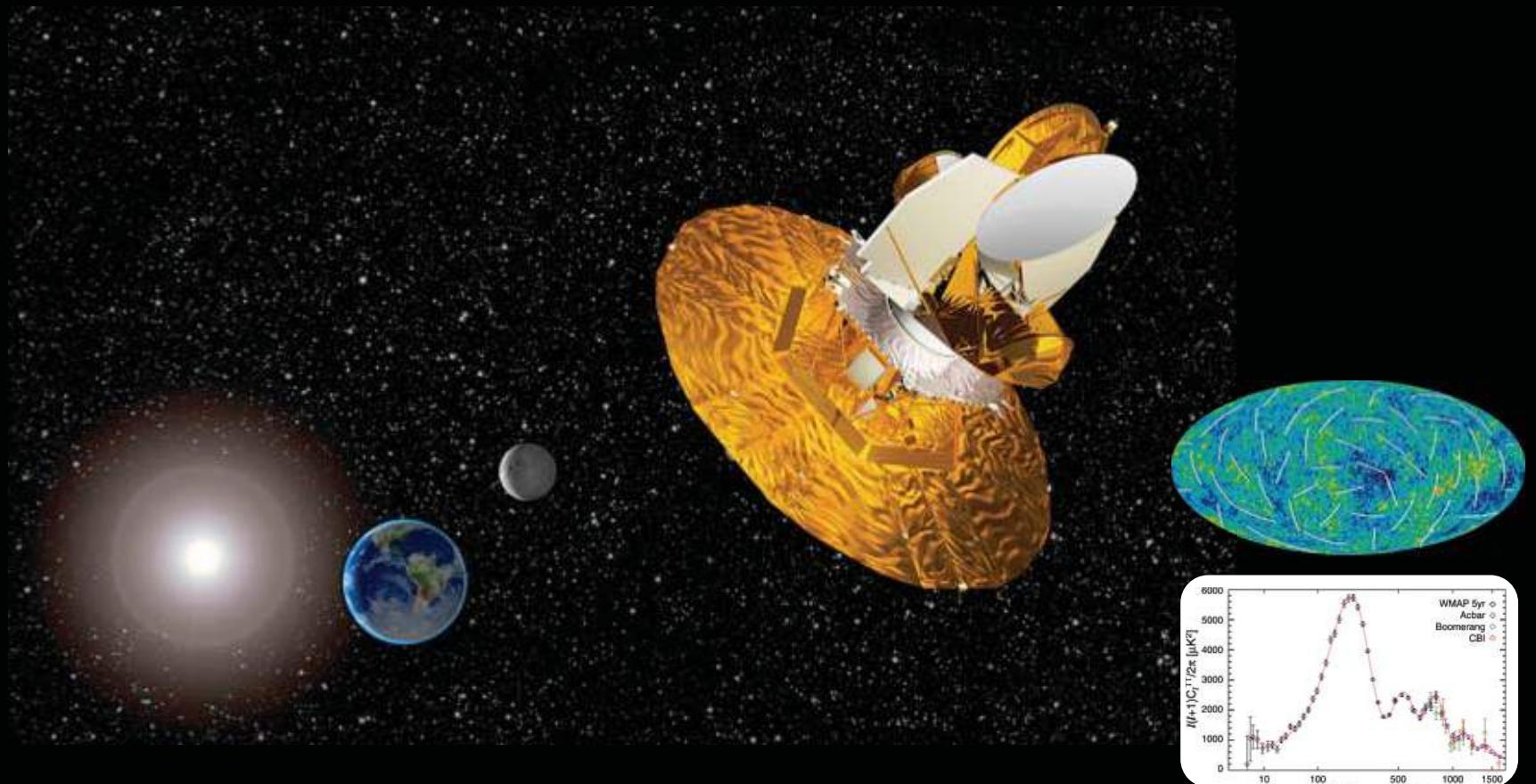


COBE 1992  
Nobel 2006 (G. Smoot & J. Mather)





Boomerang 1998-2001  
Première observation du premier pic



Wilkinson Microwave Anisotropie Probe (WMAP)  
2003-2010  
Observation des trois premiers pics  
Informations sur la polarisation



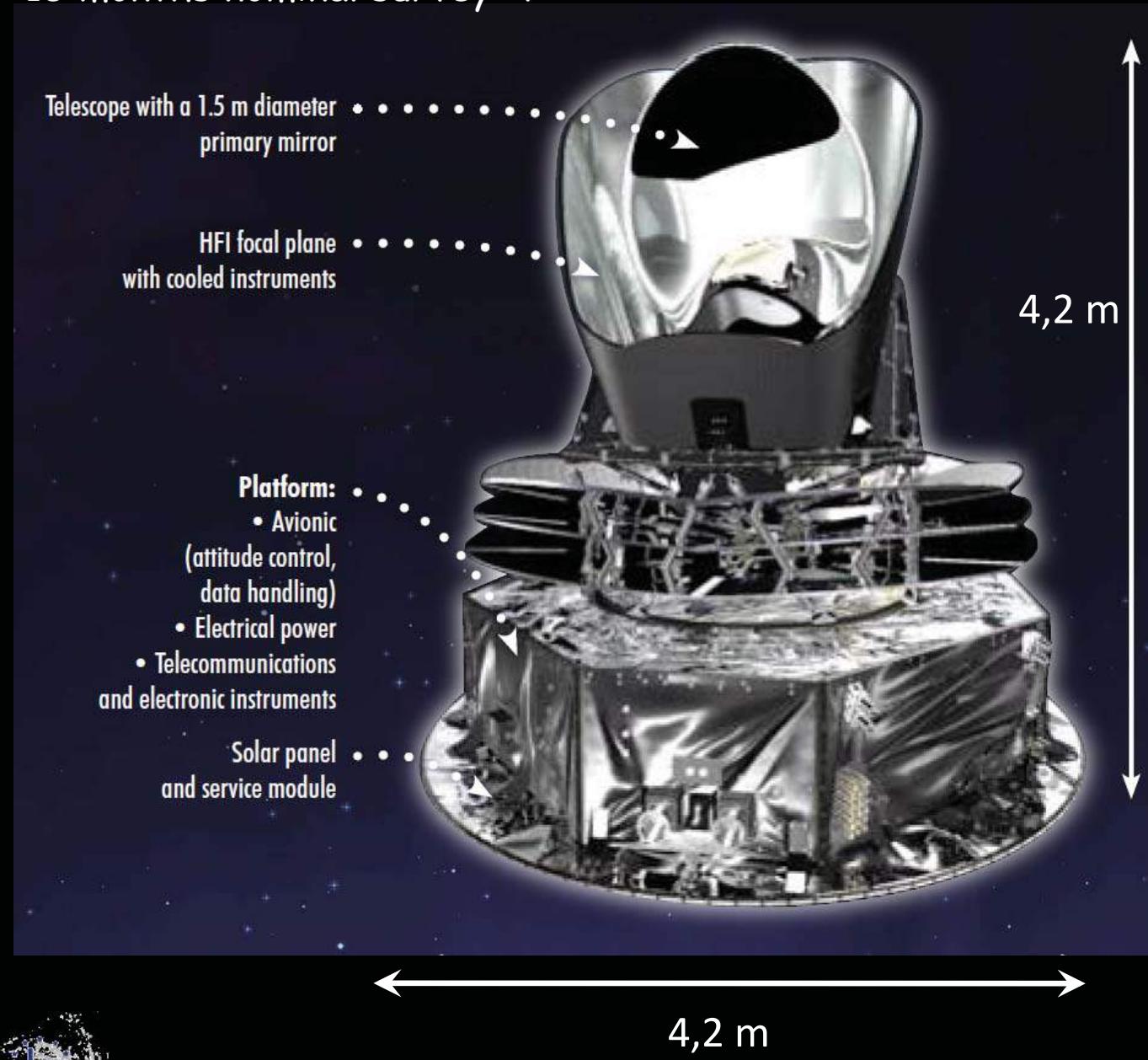
**PLANCK**

2000 Kg

1600 W consumption

2 instruments - HFI & LFI

15 months nominal survey+4



50 000 electronic components

36 000 l  $^4\text{He}$

12 000 l  $^3\text{He}$

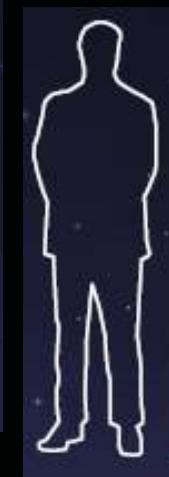
11 400 documents

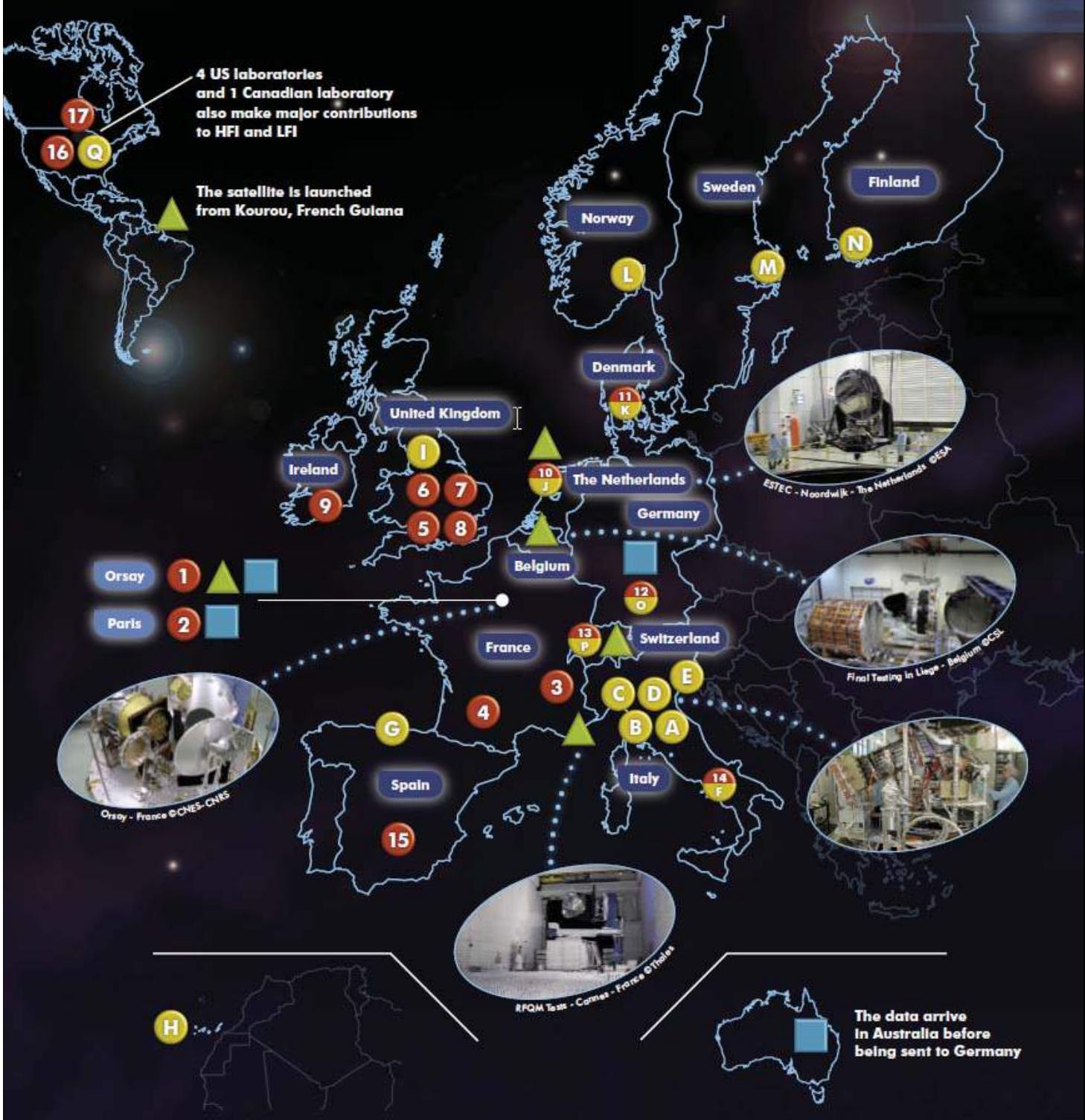
20 years between the first project and first results (2013)

6c per European per year

16 countries

400 researchers among 1000





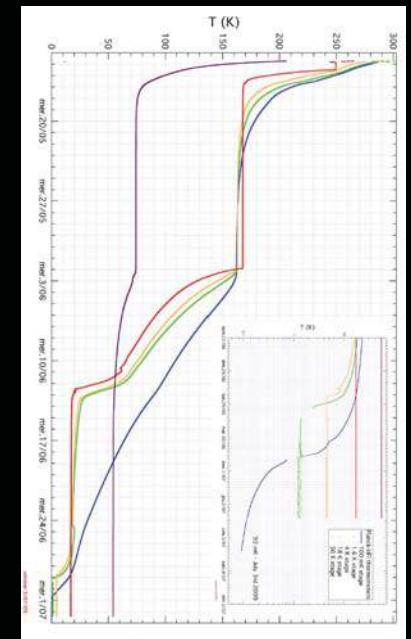
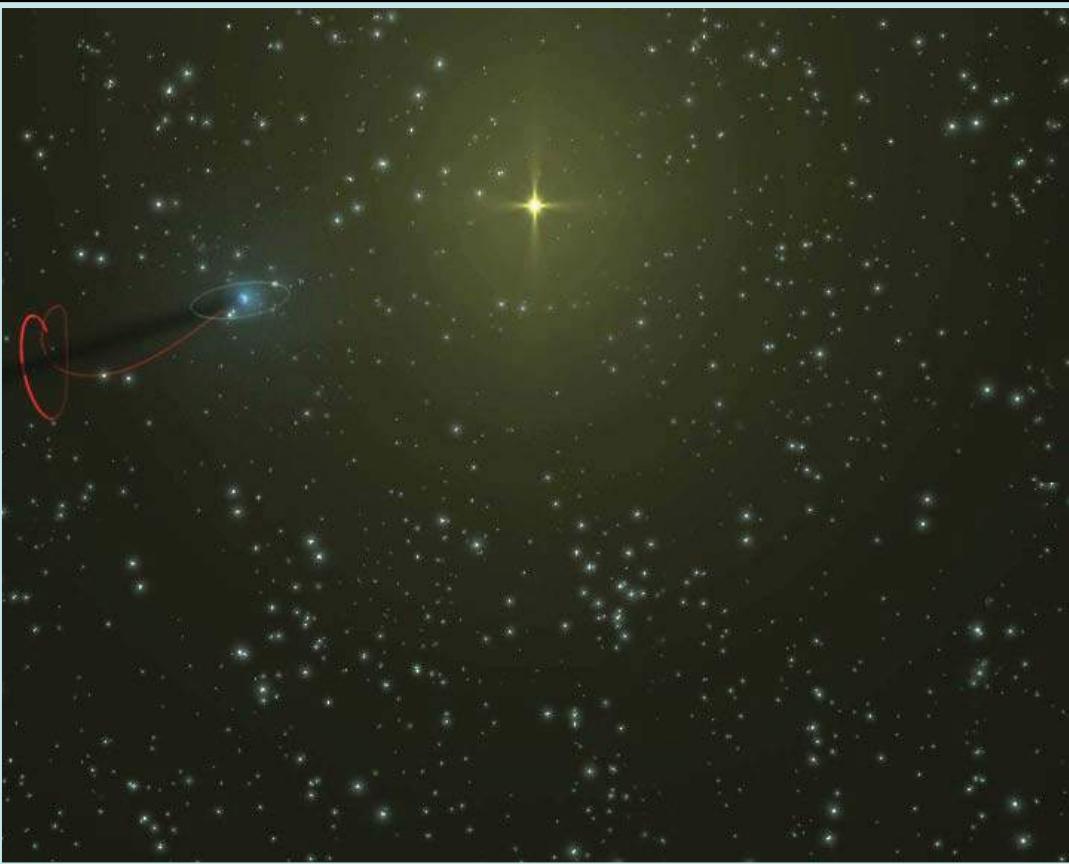
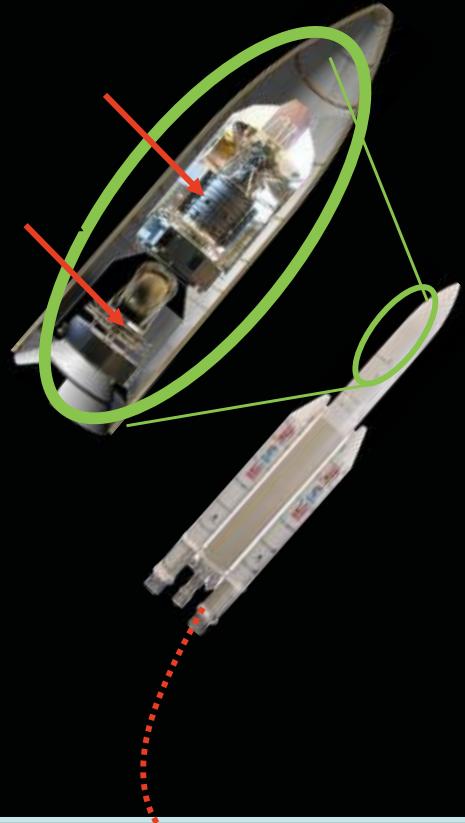
#### Research Laboratories in the HFI Collaboration

- 1 Institut d'Astrophysique Spatiale, Orsay (F)
- 1 Laboratoire de l'Accélérateur Linéaire, Orsay (F)
- 1 Commissariat à l'Energie Atomique, Gif-sur-Yvette (F)
- 2 Institut d'Astrophysique de Paris, Paris (F)
- 2 Laboratoire d'Étude du Rayonnement et de la Matière en Astrophysique, Paris, (F)
- 2 AstroParticule et Cosmologie, Paris (F)
- 3 Laboratoire de Physique Subatomique et de Cosmologie, Grenoble (F)
- 3 Institut Louis Néel, Grenoble (F)
- 4 Centre d'Etudes Spatiales des Rayonnements, Toulouse (F)
- 5 Cardiff University, Cardiff (UK)
- 6 Rutherford Appleton Laboratory, Chilton (UK)
- 7 Institute of Astronomy, Cambridge (UK)
- 7 Mullard Radio Astronomy Observatory, Cambridge (UK)
- 8 Imperial College, London (UK)
- 9 National University of Ireland, Maynooth (IR)
- 10 Space Science Dpt of ESA, Noordwijk (NL)
- 11 Danish Space Research Institute, Copenhagen (DK)
- 12 Max-Planck-Institut fuer Astrophysik, Garching (D)
- 13 Université de Genève , Geneva (CH)
- 14 University La Sapienza, Rome (I)
- 15 Universidad de Granada, Granada (E)
- 16 California Institute of Technology, Pasadena (USA)
- 16 Jet Propulsion Laboratory, Pasadena (USA)
- 16 Stanford University, Stanford (USA)
- 17 Canadian Institute for Theoretical Astrophysics, Toronto (Canada)

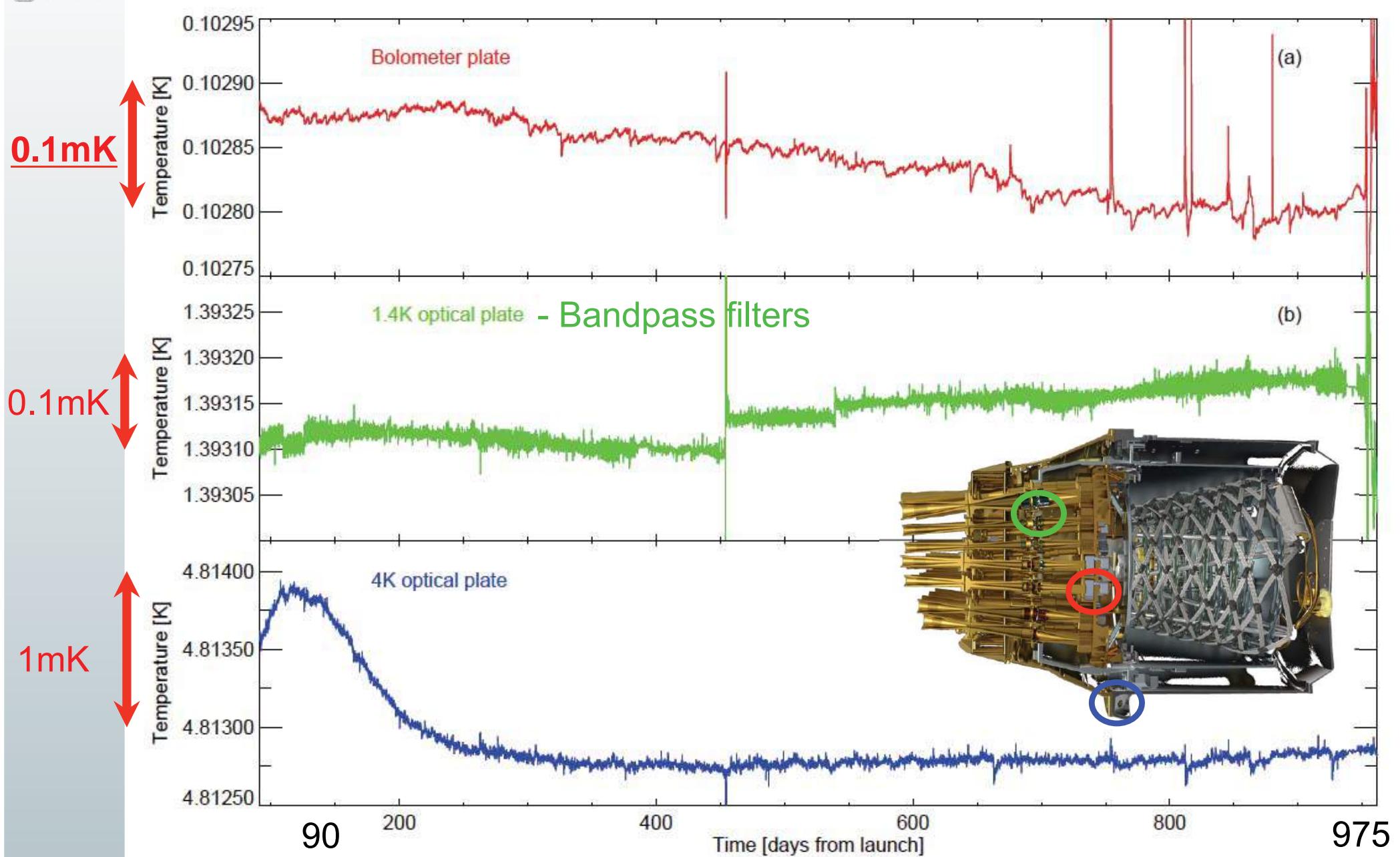
#### Research Laboratories in the LFI Collaboration

- A Istituto Nazionale di Astrofisica Spaziale et Fisica Cosmica, Bologna (I)
- B Istituto CAISM, Firenze (I)
- C Istituto IASF (CNR), Milano (I)
- C Istituto di Fisica del Plasma IFP (CNR), Milano (I)
- D Osservatorio Astronomico di Padova, Padova (I)
- E Osservatorio Astronomico di Trieste, Trieste (I)
- E SISSA, Trieste (I)
- F Istituto IFSI, Roma (I)
- F Università Tor Vergata, Roma (I)
- G Istituto de Fisica de Cantabria, Santander (E)
- H Instituto de Astrofísica de Canarias, La Laguna (E)
- I Jodrell Bank Observatory, Macclesfield (UK)
- J Space Science Dpt of ESA , Noordwijk (NL)
- K Danish Space Research Institute , Copenhagen (DK)
- K Theoretical Astrophysics Center, Copenhagen (DK)
- L University of Oslo, Oslo (N)
- M Chalmers University of Technology, Göteborg (S)
- N Millimetre Wave Laboratory, Espoo (F)
- O Max-Planck-Institut fuer Astrophysik, Garching (D)
- P Université de Genève, Geneva (CH)
- Q University of California (Berkeley), Berkeley (USA)
- Q University of California (Santa Barbara), Santa Barbara (USA)
- Q Jet Propulsion Laboratory, Pasadena (USA)



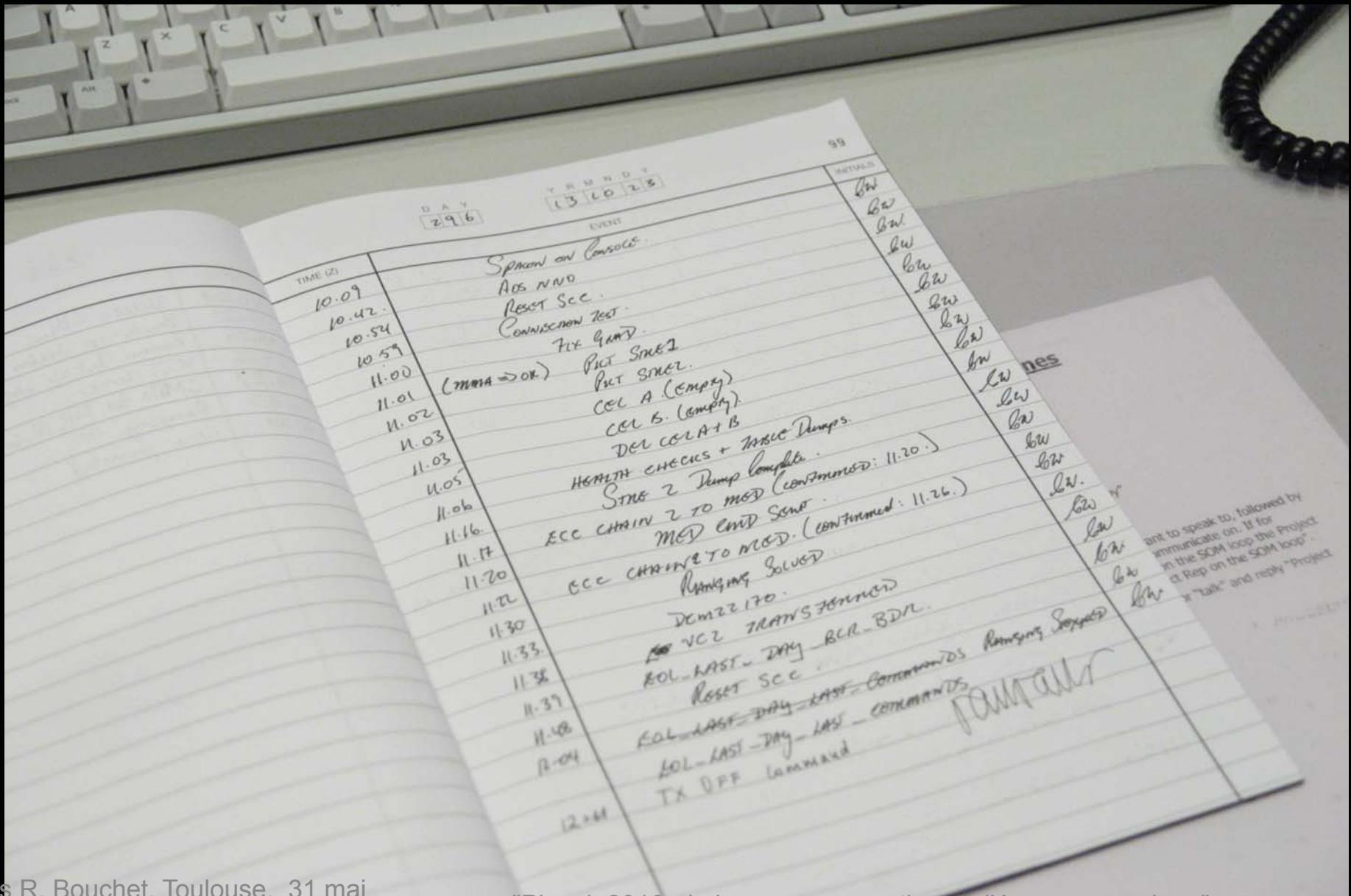


# When I say very stable...



# Last commands

## Darmstadt, October 23<sup>rd</sup> 2013



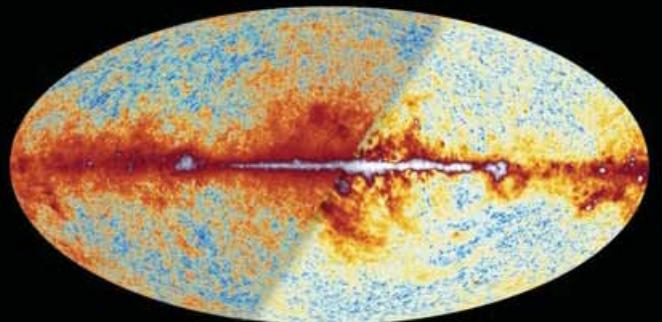
Want to speak to, followed by  
immaculate on. If for  
in the SOM loop. Project  
et Rep on the SOM loop.  
or "talk" and reply "Project

# A la Cave de l'IAP...

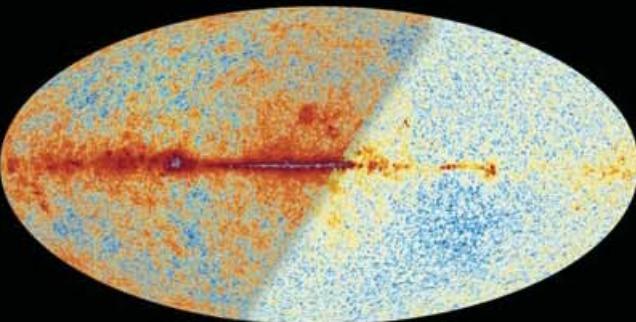
+ CC/CINECA/  
Darwin/NERSC...



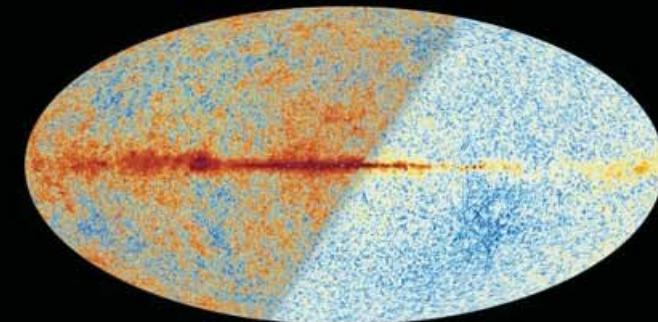
# Disponible maintenant, chez vous!



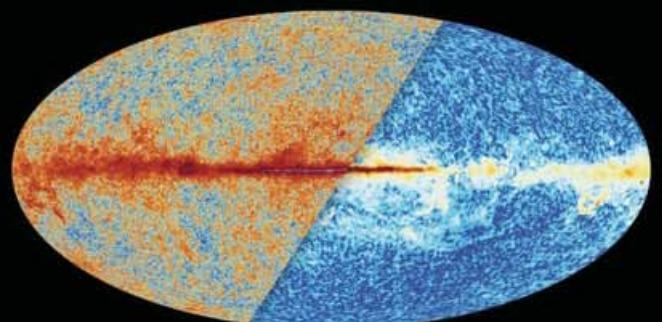
30 GHz



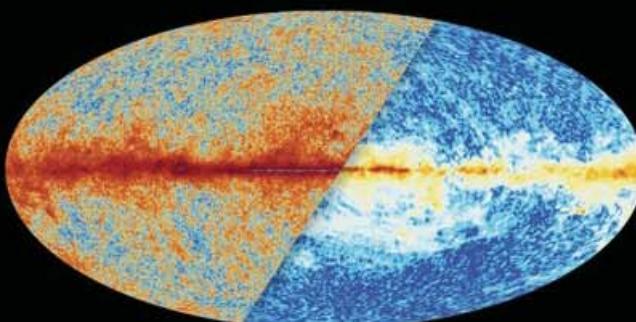
44 GHz



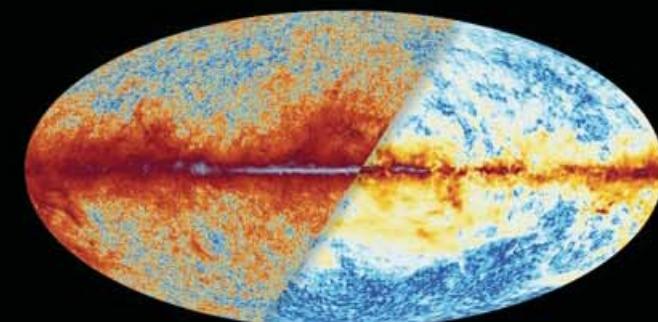
$3.5\mu\text{K.deg}, 13'$  70 GHz



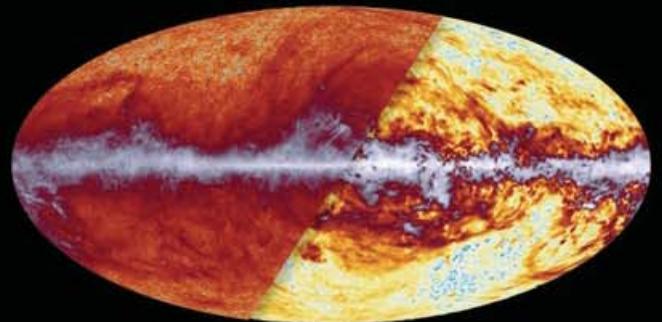
$1.3\mu\text{K.deg}, 9.7'$  100 GHz



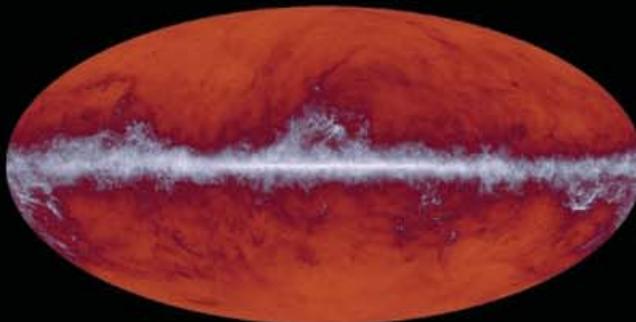
$0.5\mu\text{K.deg}, 7.3'$  143 GHz



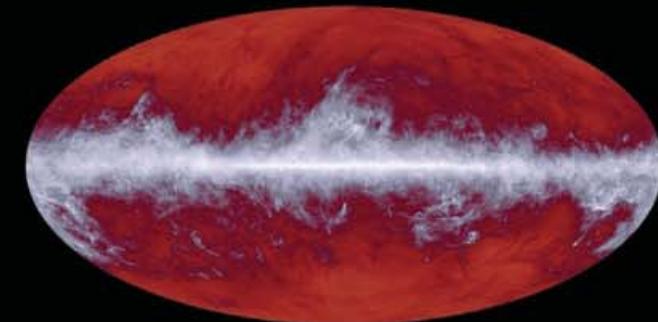
$0.8\mu\text{K.deg}, 5.0'$  217 GHz



353 GHz

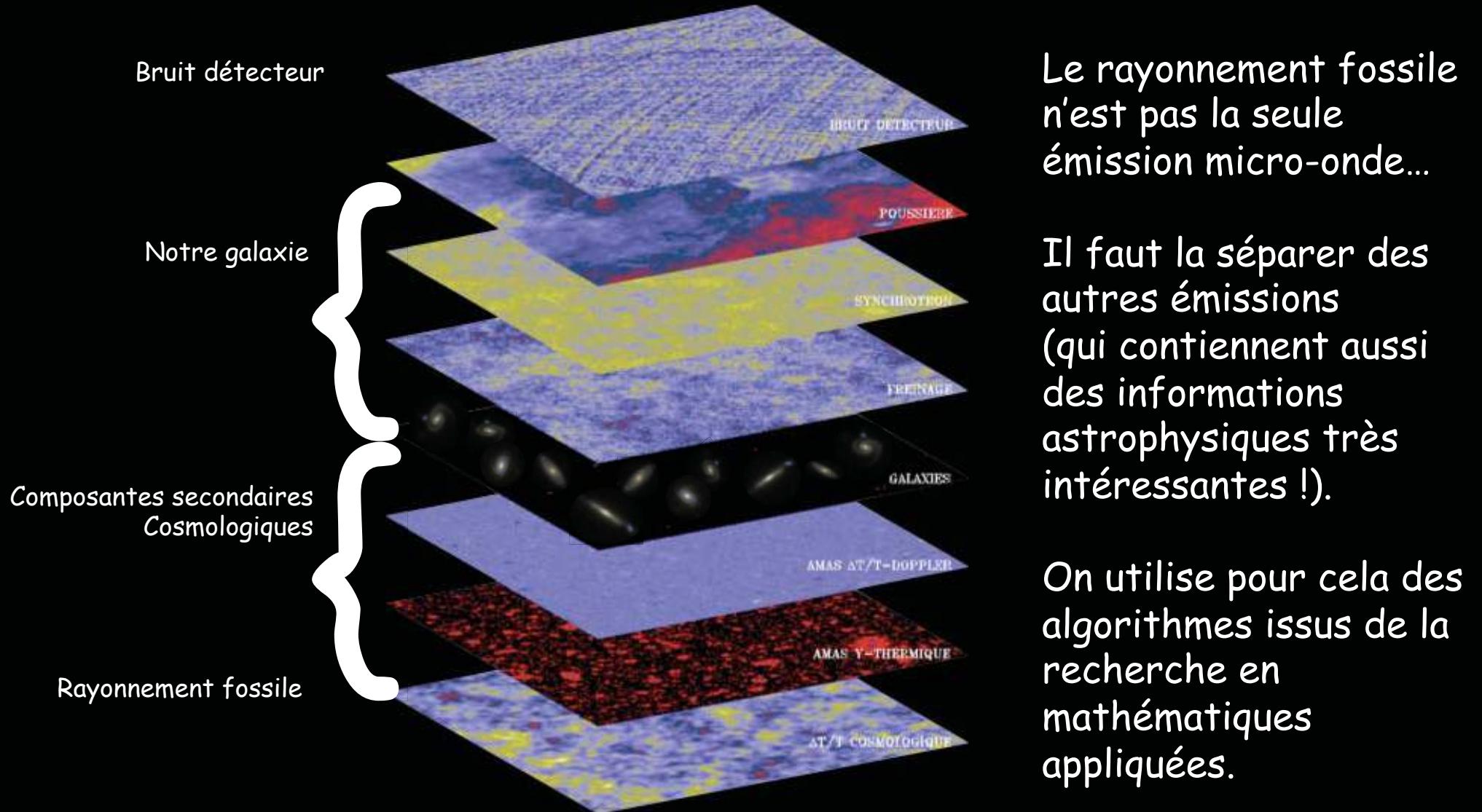


545 GHz



857 GHz

# Un autre défi... Séparer les différentes composantes

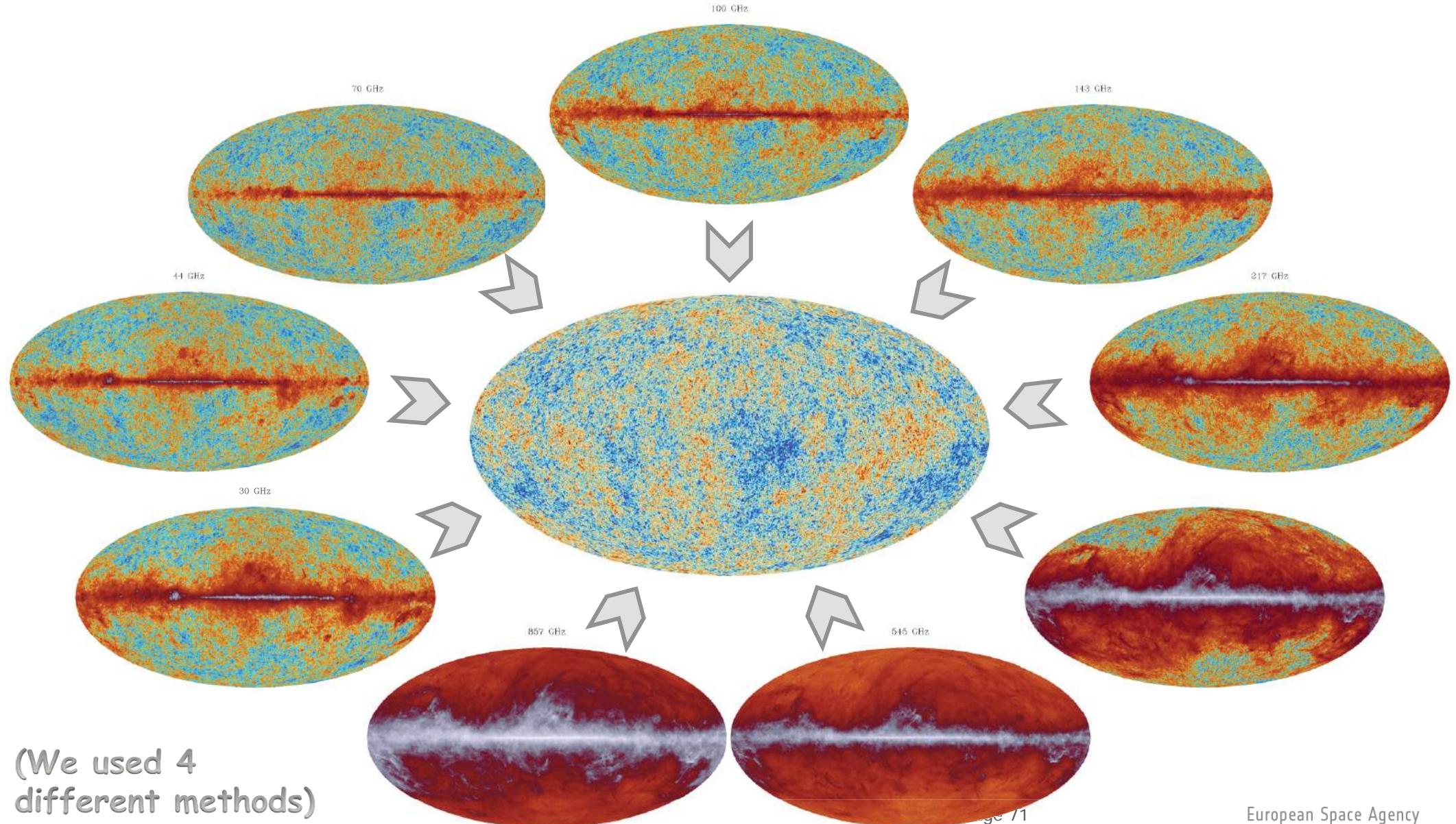


# C'est le problème du cocktail...

Quand il faut isoler une conversation à l'aide de divers enregistrements



# Cleaning the background from its 7 veils



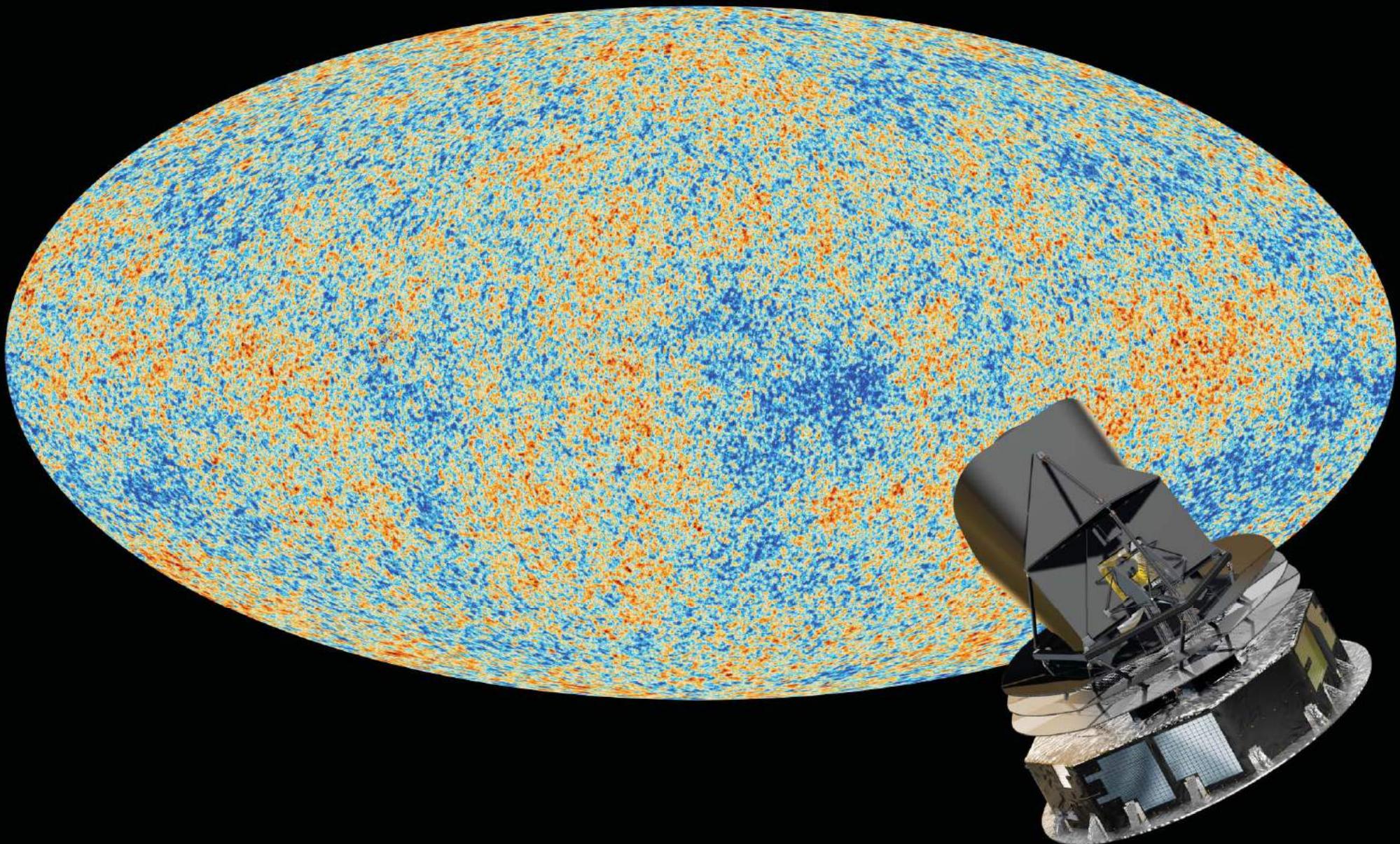
(We used 4  
different methods)

"Planck 2016: de la mousse quantique à l'éponge cosmique"

European Space Agency

Page 71  
3% of the CMB sky replaced by a Gaussian Random realisation

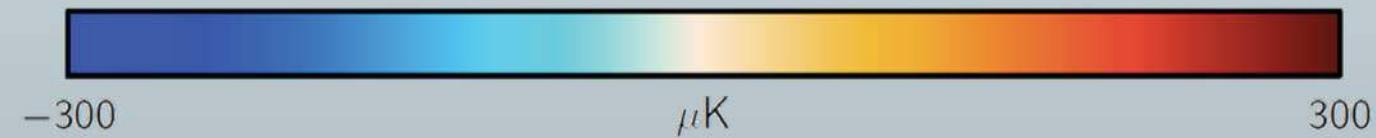
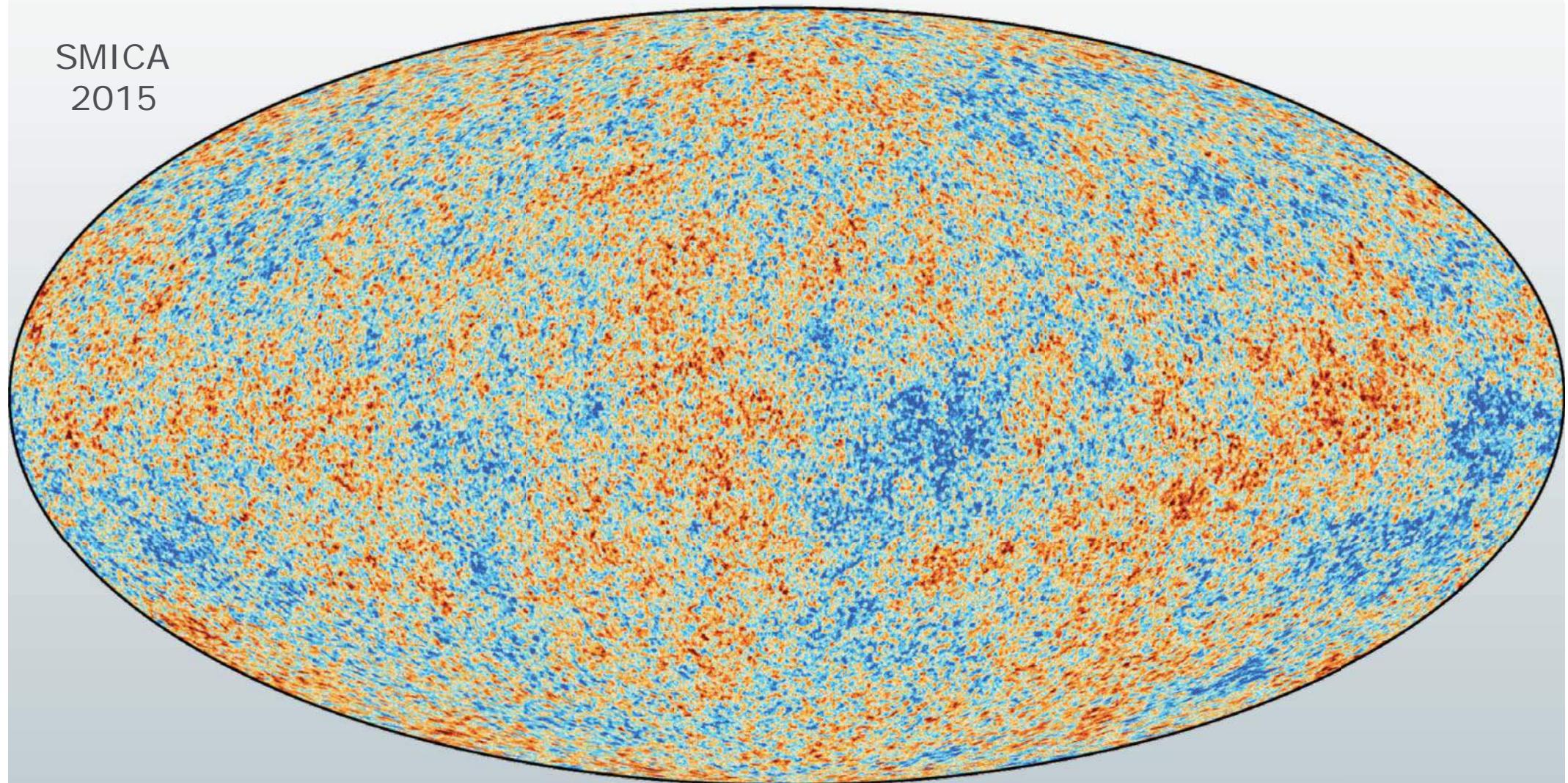
# La carte exhaustive des anisotropies de température relevée par Planck





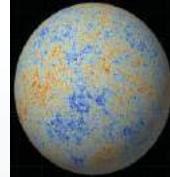
# Planck 2015 T anisotropies map

SMICA  
2015





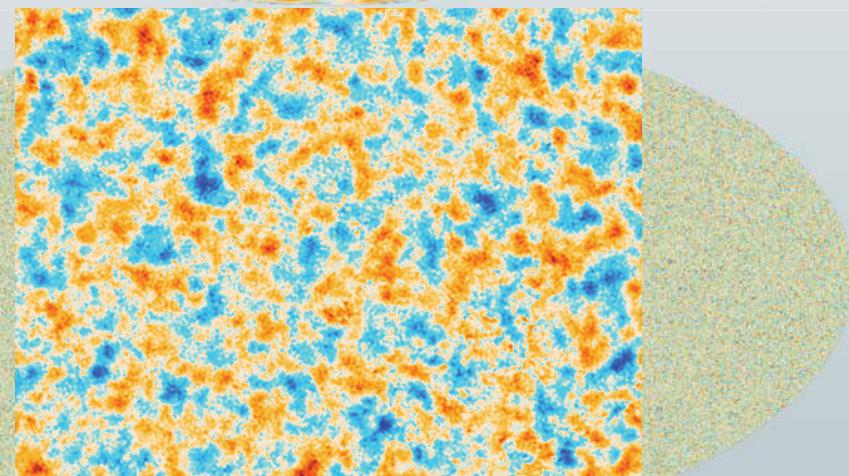
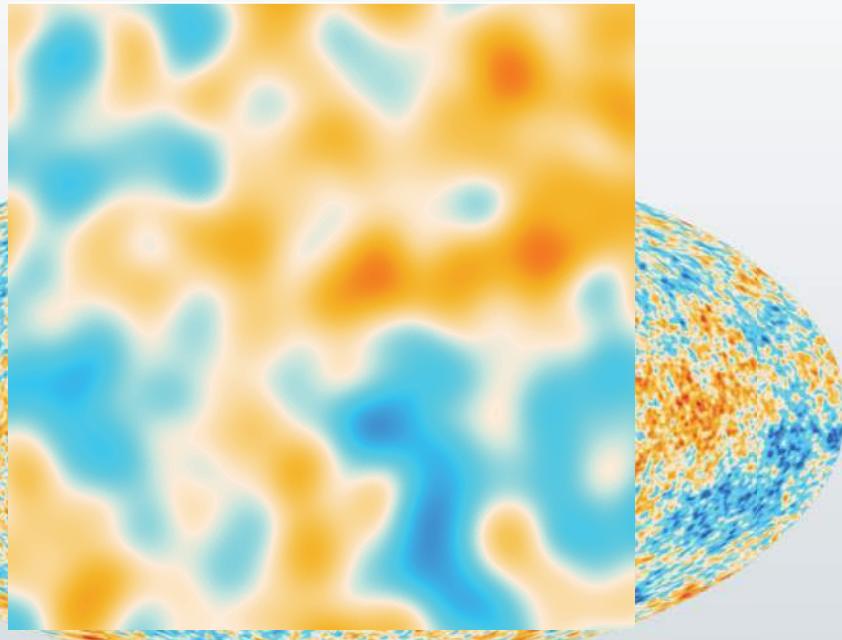
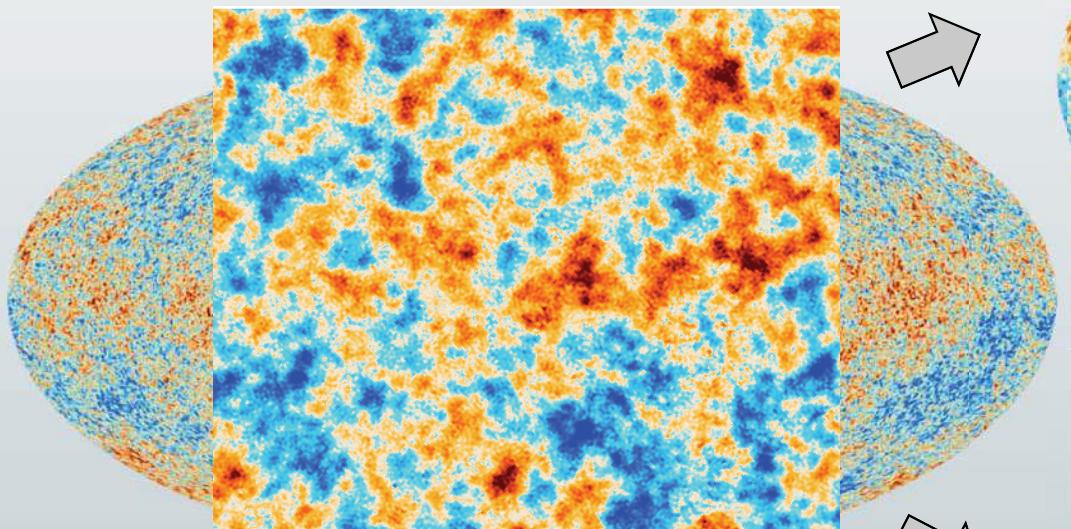
# Une fenêtre unique...



Smoothed map (suppressing scales  $\theta < 1$  deg) :

Quantum Fluctuations imprinted

When the age of the Universe was in the  
interval  $[10^{-39}, 10^{-12}]$  seconds



Difference map (scales  $\theta < 1$  deg) :

Acoustic oscillations at small scales

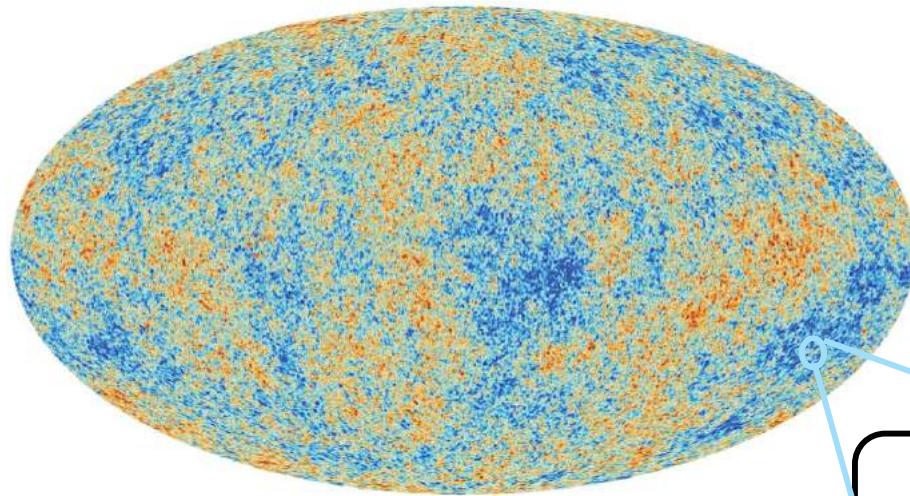
$< ct$  when  $t=380\ 000$  years ( $\sim 150$  Mpc today).

Which allows to take a census of the Universe content

# Ce qu'en dit la théorie...

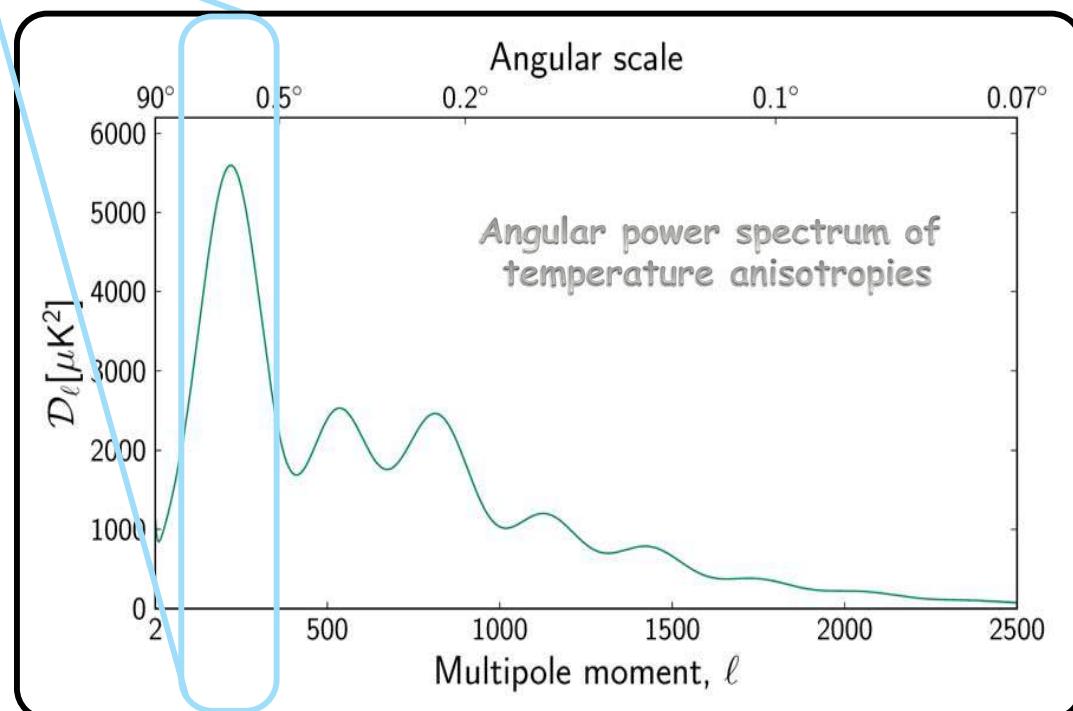


## (bien avant les observations...)

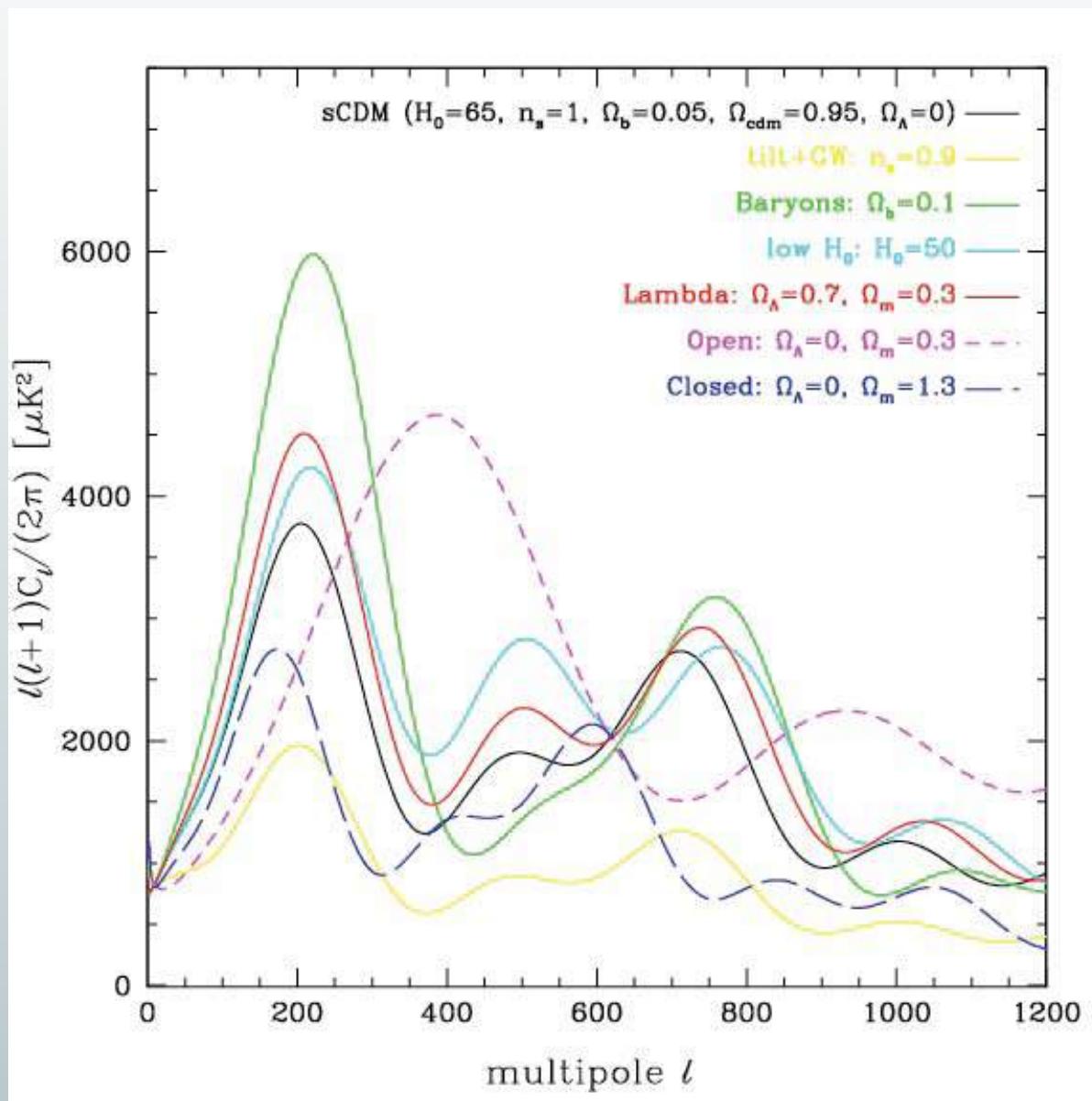


On ne peut prédire la carte des anisotropies, telle que nous l'observons...

Mais on peut prédire ses propriétés statistiques !  
(comme par exemple la hauteur typique des vagues en fonction de leur distance crête à crête)

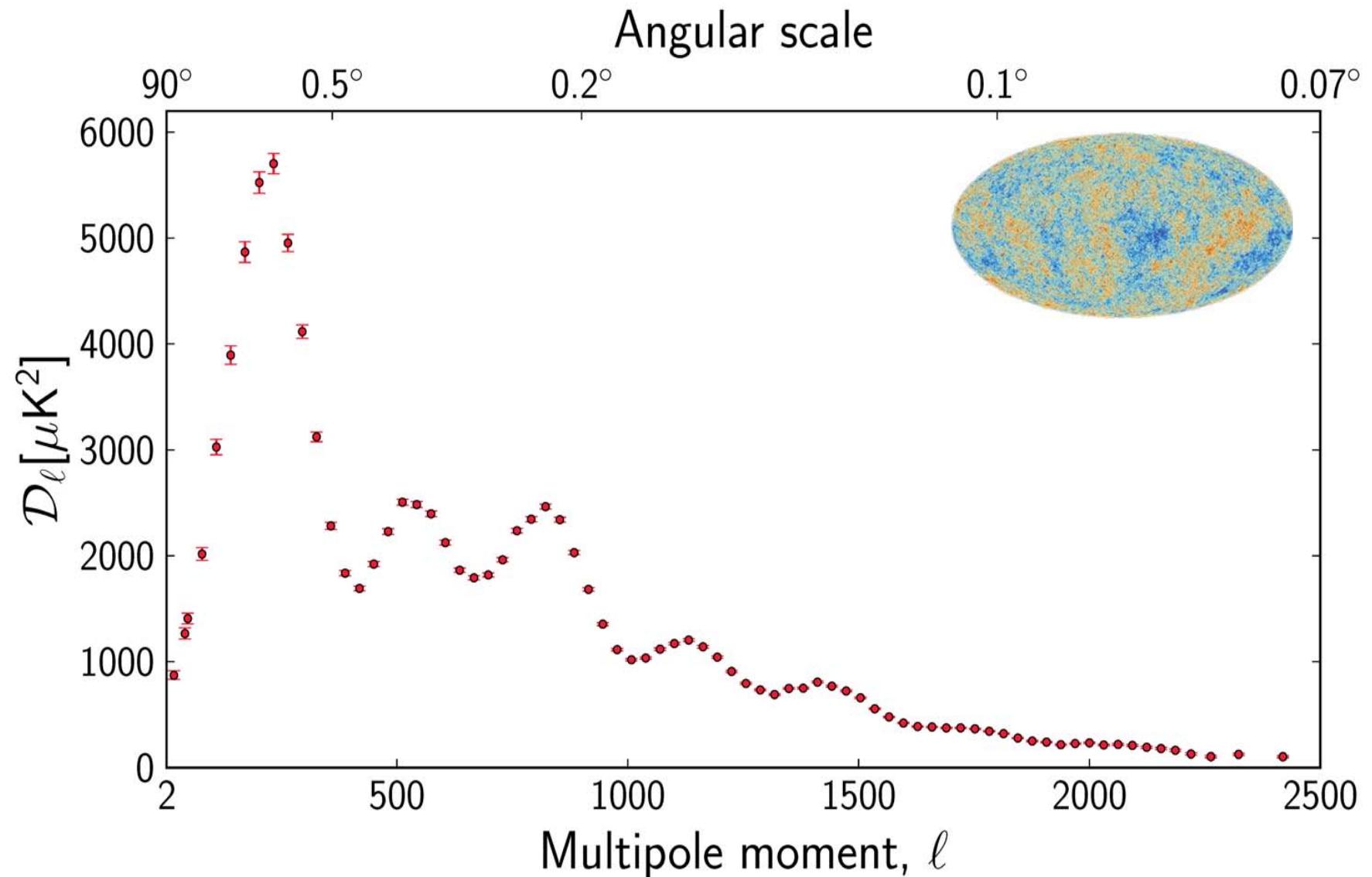


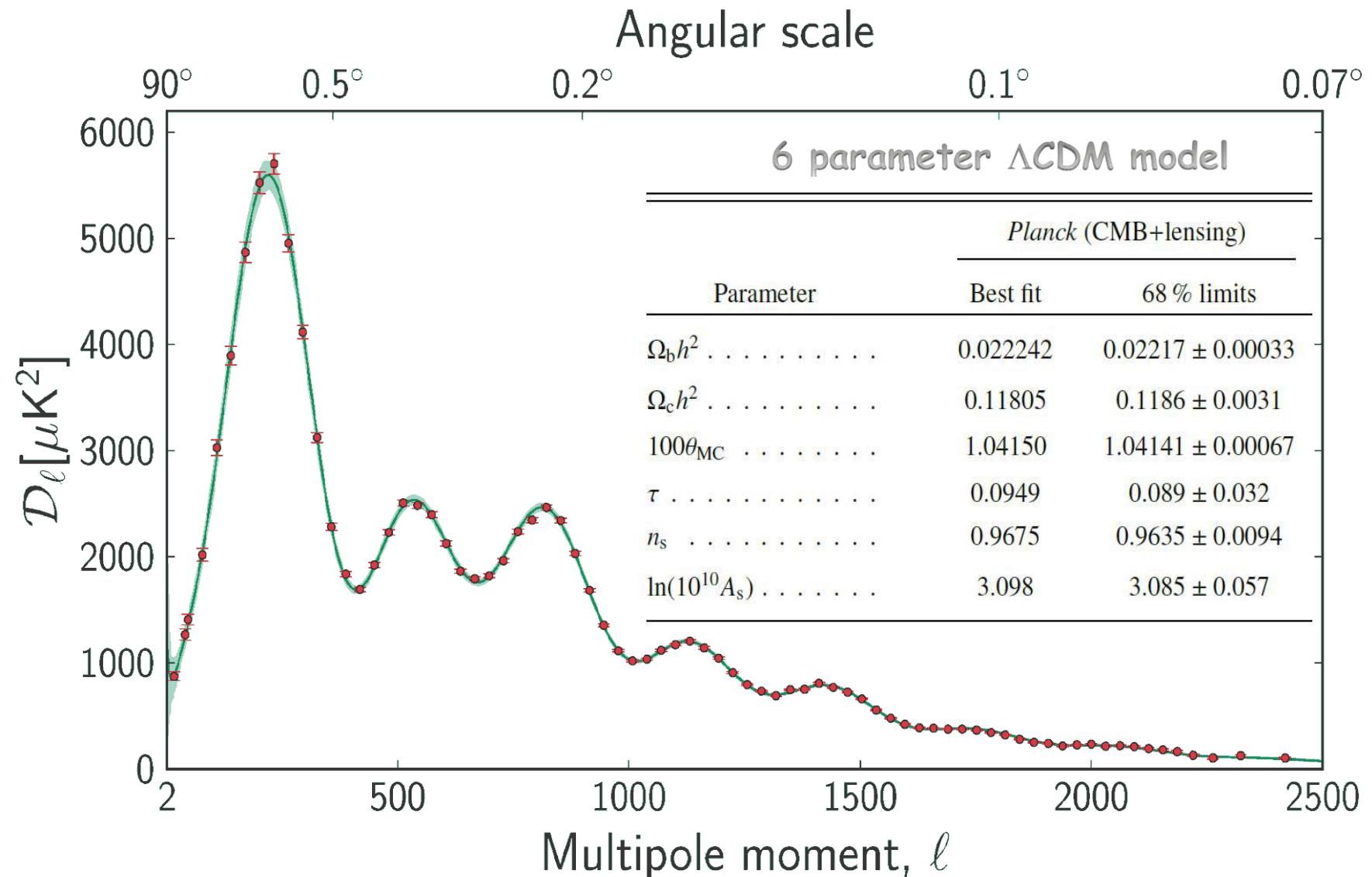
# Cosmic imprints



Peebles, Yu, Sachs & Wolf, Sunyaev, Zeldovich, Silk, Vittorio, Wilson, Mukhanov, Chibisov, Bardeen, Linde, Bond, Efstathiou, Bouchet, Bennett, Gott, Kaiser, Stebbins, Allen, Shellard, Seljak, Zaldarriaga, Kamionkowski, Hu, ...

# The Planck power spectrum of Temperature anisotropies





# Base $\Lambda$ CDM model with 6 parameters

3 parameters to set (through General Relativity) the dynamics of the universe,

1 parameter to capture the effect of reionisation (end of the dark ages),

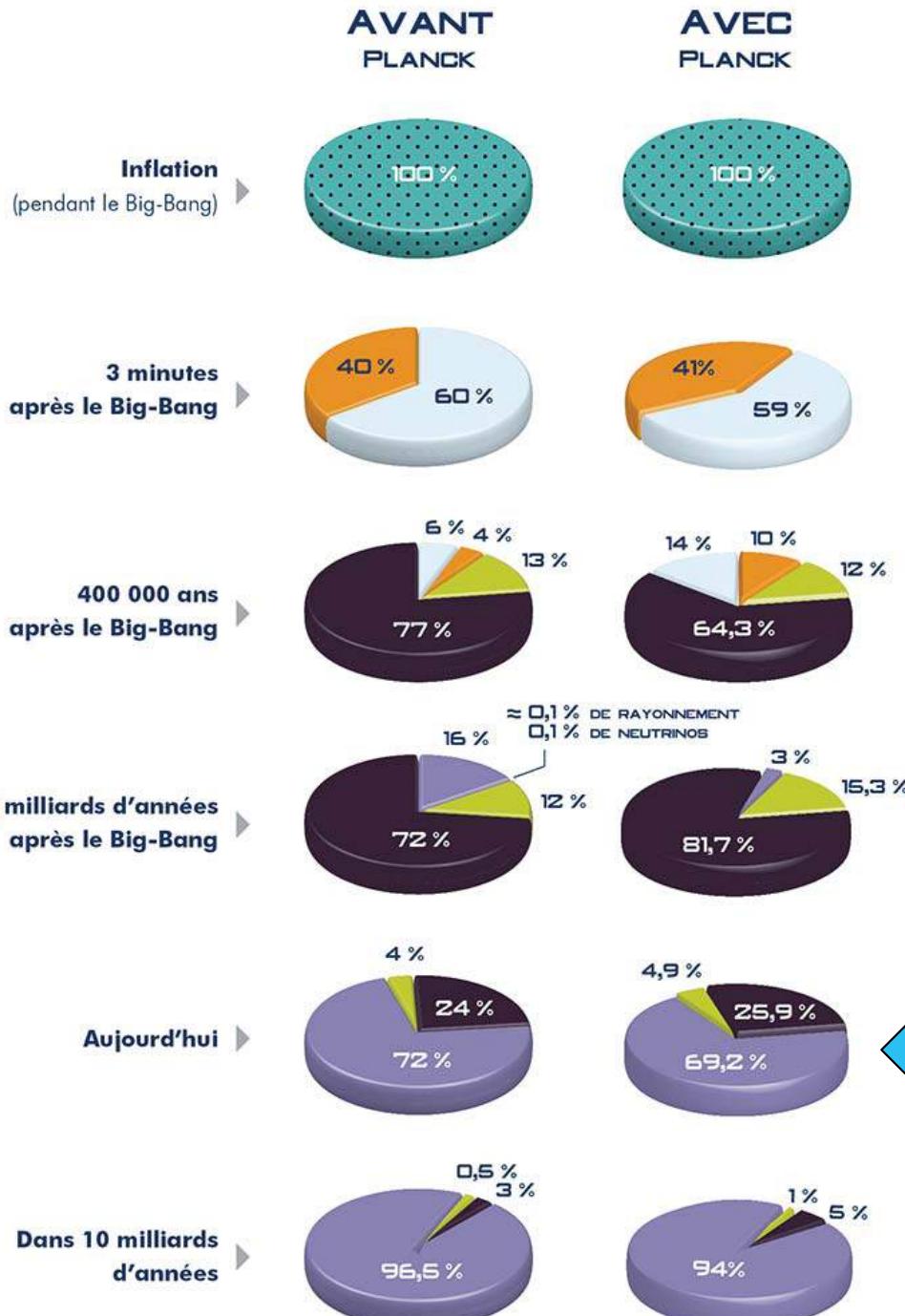
2 parameters to describe the primordial fluctuations.

Flat spatial geometry.

- $\Omega_b h^2$  Baryon density today - The amount of ordinary matter
- $\Omega_c h^2$  Cold dark matter density today - only weakly interacting
- $\Theta$  Sound horizon size when optical depth  $\tau$  reaches unity  
*(Distance traveled by a sound wave since inflation, when universe became transparent at recombination at  $t \sim 380\,000$  years)*
- $\tau$  Optical depth at reionisation (due to Thomson scattering of photons on  $e^-$ ), i.e. fraction of the CMB photons re-scattered during that process
- $A_s$  Amplitude of the curvature power spectrum  
*(Overall contrast of primordial fluctuations)*
- $n_s$  Scalar power spectrum power law index  
*( $n_s - 1$  measures departure from scale invariance)*
- Others are derived parameters within the model, in particular
  - $\Omega$  "Dark Energy" fraction of the critical density (derived only if assumed flat)
  - $H_0$  the expansion rate today (in km/s per Mpc of separation)
  - $t_0$  the age of the universe (in Gy)

# Data compression...

- ~ 900 billion samples in ~100 Timelines
- 9 maps of ~50 million pixels each (23 with Polar.)
- 1 CMB map of ~50 million pixels (+P one)
- 2 million harmonic modes ( $2l+1$  m-modes per  $l$ )
- Fit with just 6 parameters
- *With no significant evidence for a 7<sup>th</sup>*



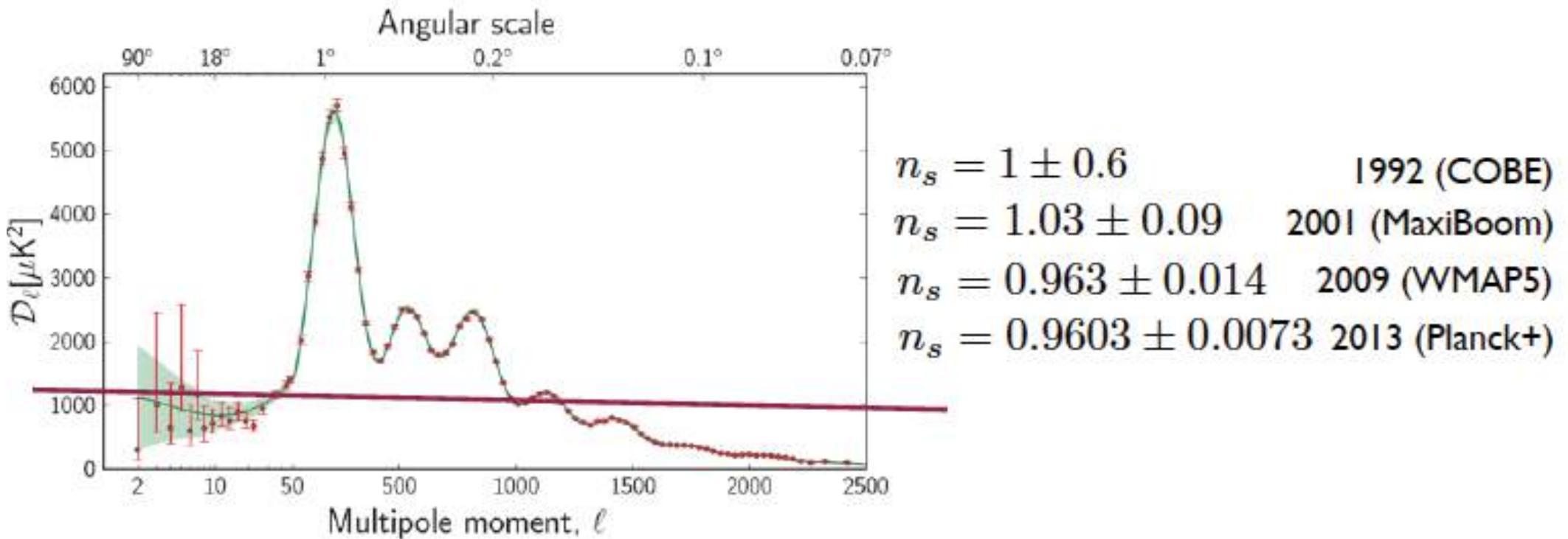
Today's composition

And  $H_0 = 67 \text{ km/s/Mpc}$   
within  $< 1 \text{ km/s/Mpc}$

# What is the value of $n_s$ ?

Initial Conditions: quasi-scale invariant

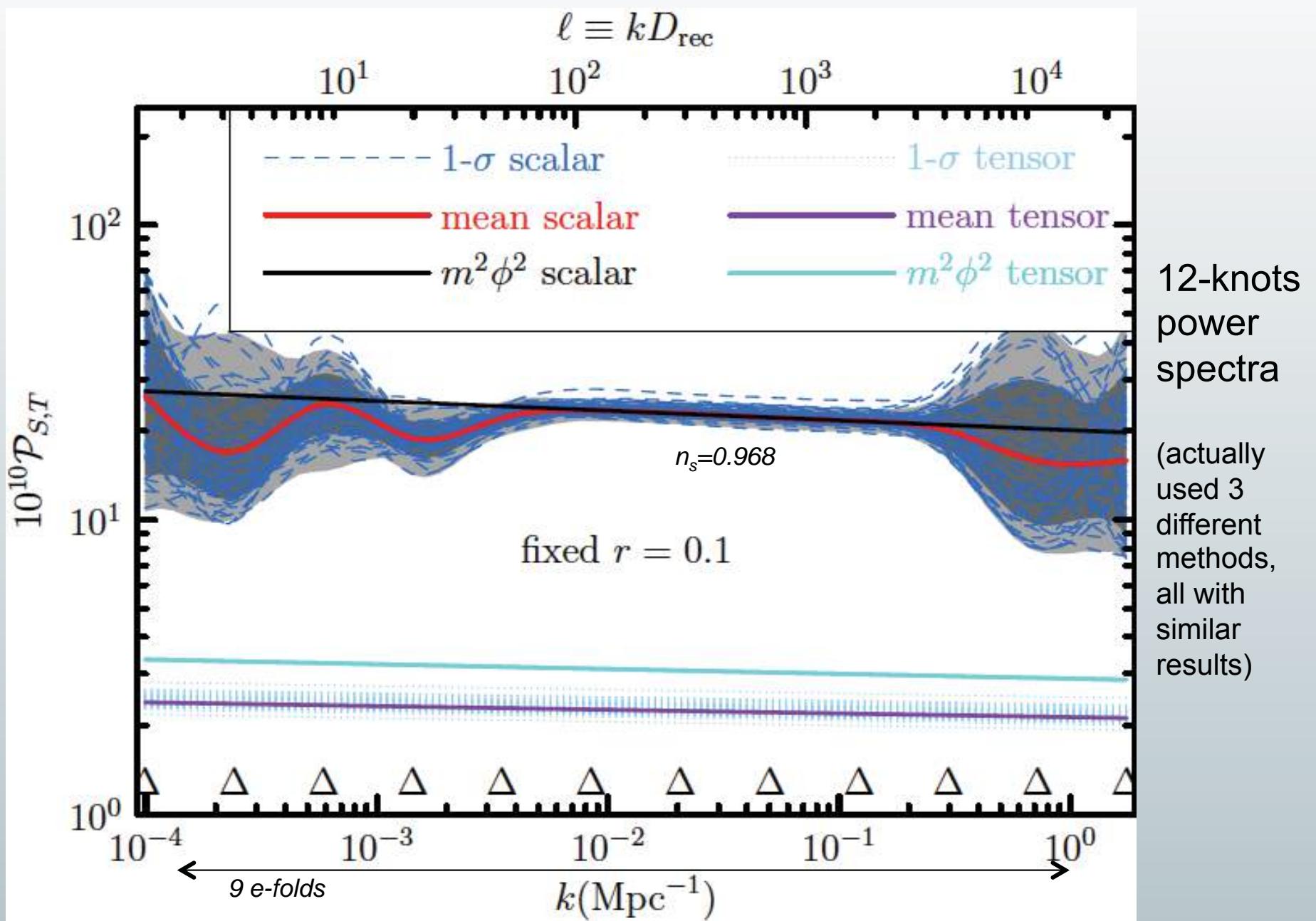
$$g_{ij} = a^2(\tau) [1 - 2\Phi] \gamma_{ij} \longrightarrow k^3 \langle |\Phi_k| \rangle \propto k^{n_s - 1}$$



Mukhanov & Chibisov (1981!): 1<sup>st</sup> calculation of (scalar) quantum fluctuation of the vacuum in expanding background.  $n_s$  must be  $\sim 0.96 < 1$  for inflation to end.

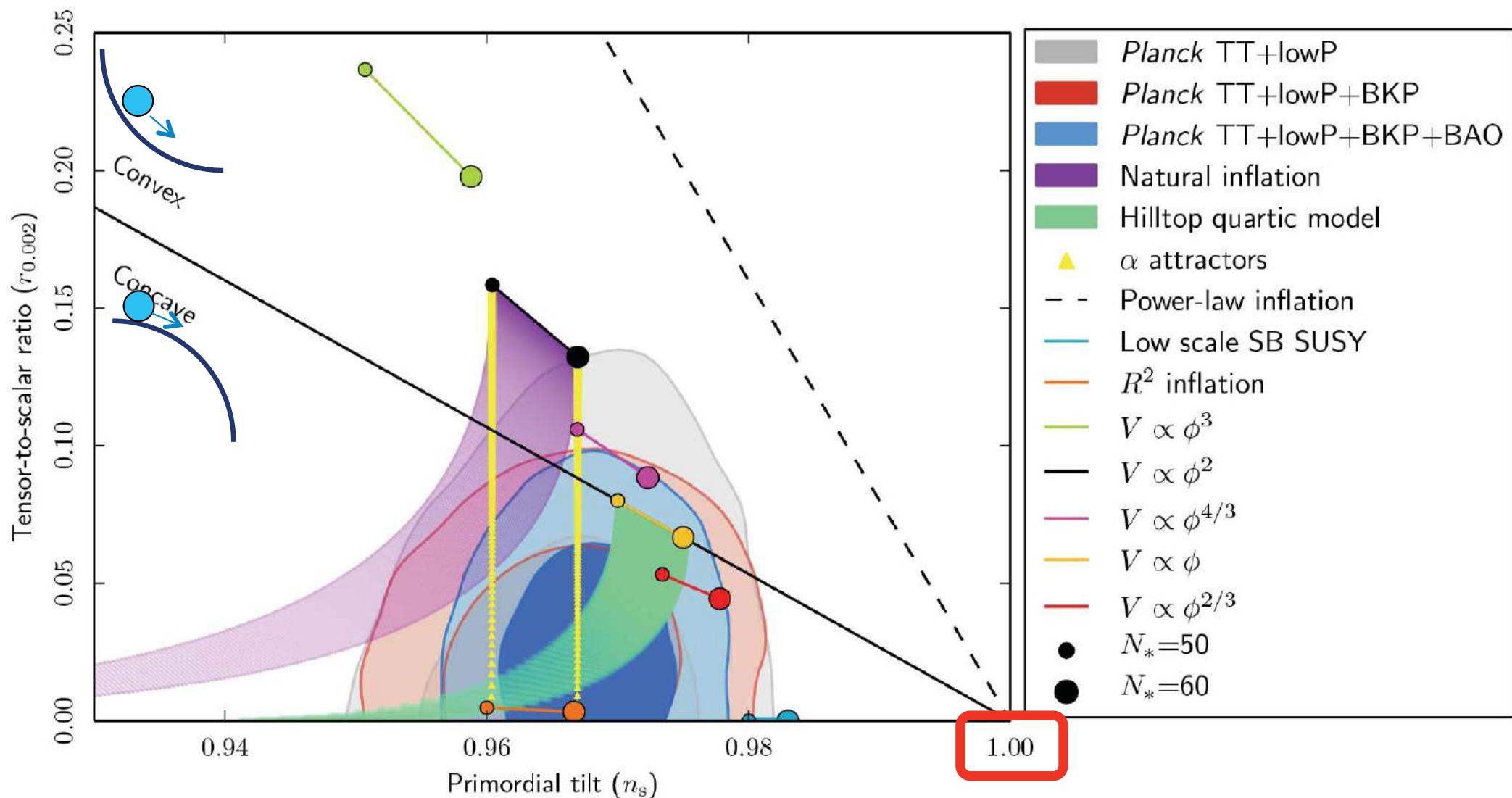
# Power spectra reconstruction

2015  
TT+lowP  
+BAO+JLA  
+Hlow



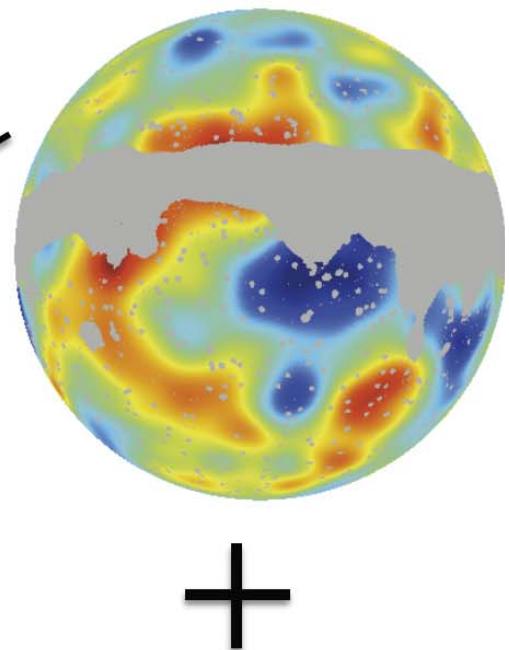
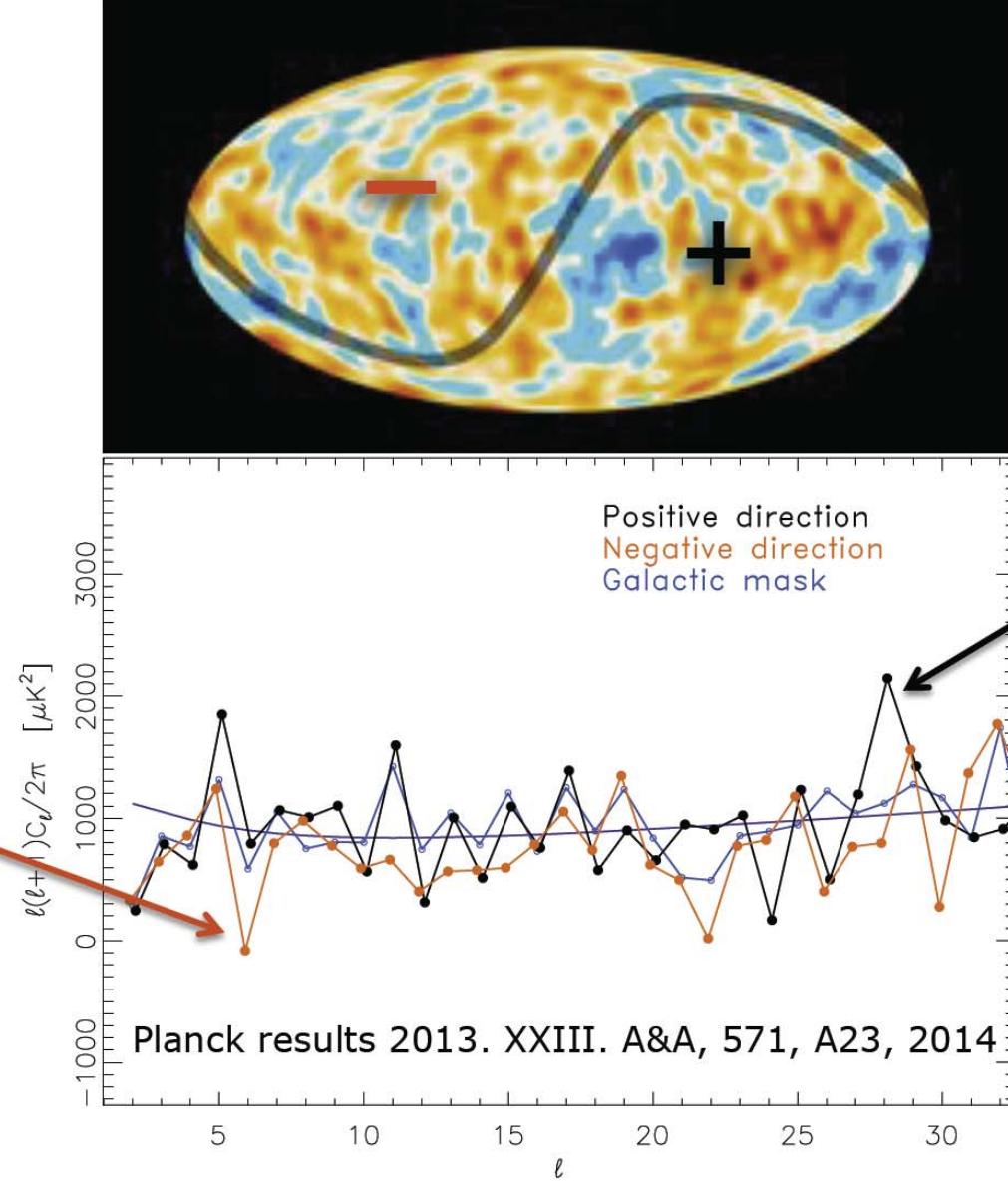
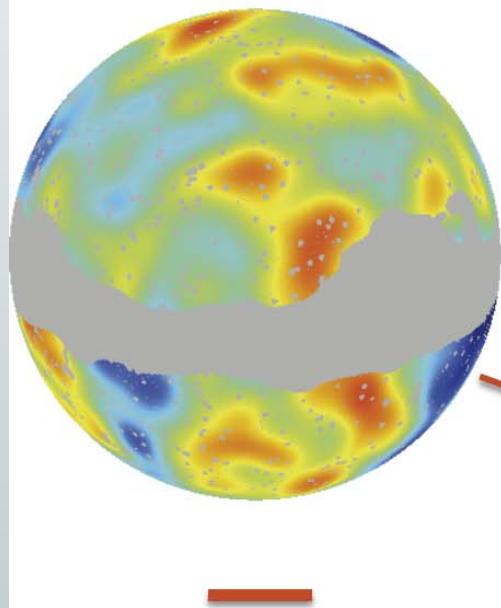
# Une clé essentielle... Planck 2015: $n_s$ vs $r$

$$V_* = (1.9 \times 10^{16} \text{ GeV})^4 (r/0.12)$$



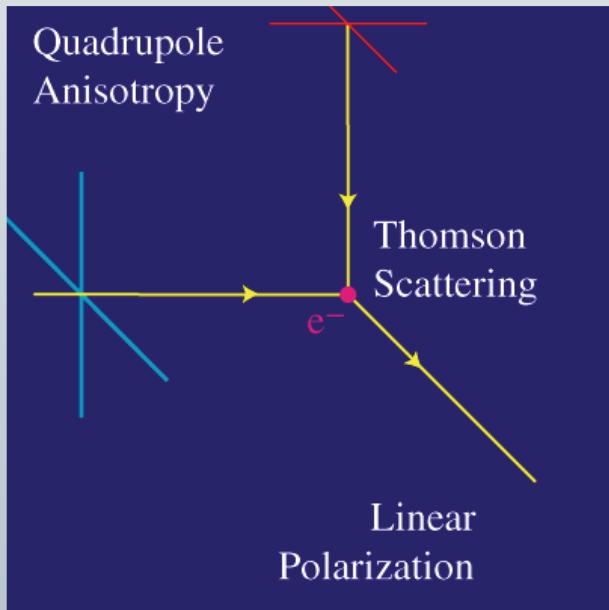
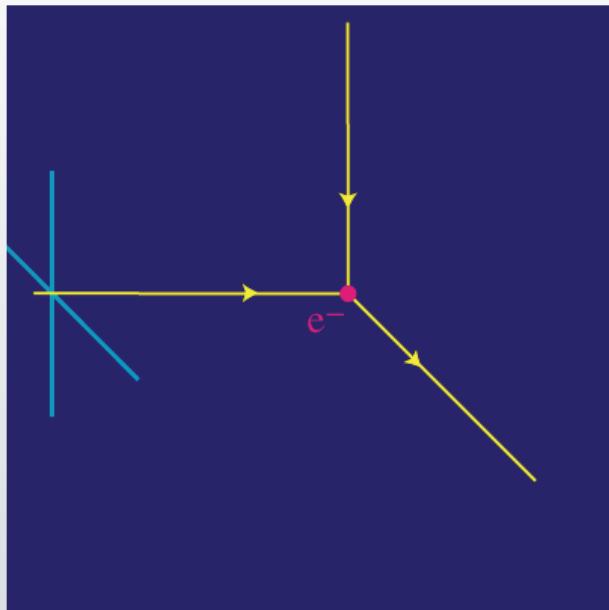
Similar (indirect)  $r$  constraint than with 2013 release ( $r_{0.002} < 0.10$  @ 95% CL vs 0.11)

# Power asymmetry

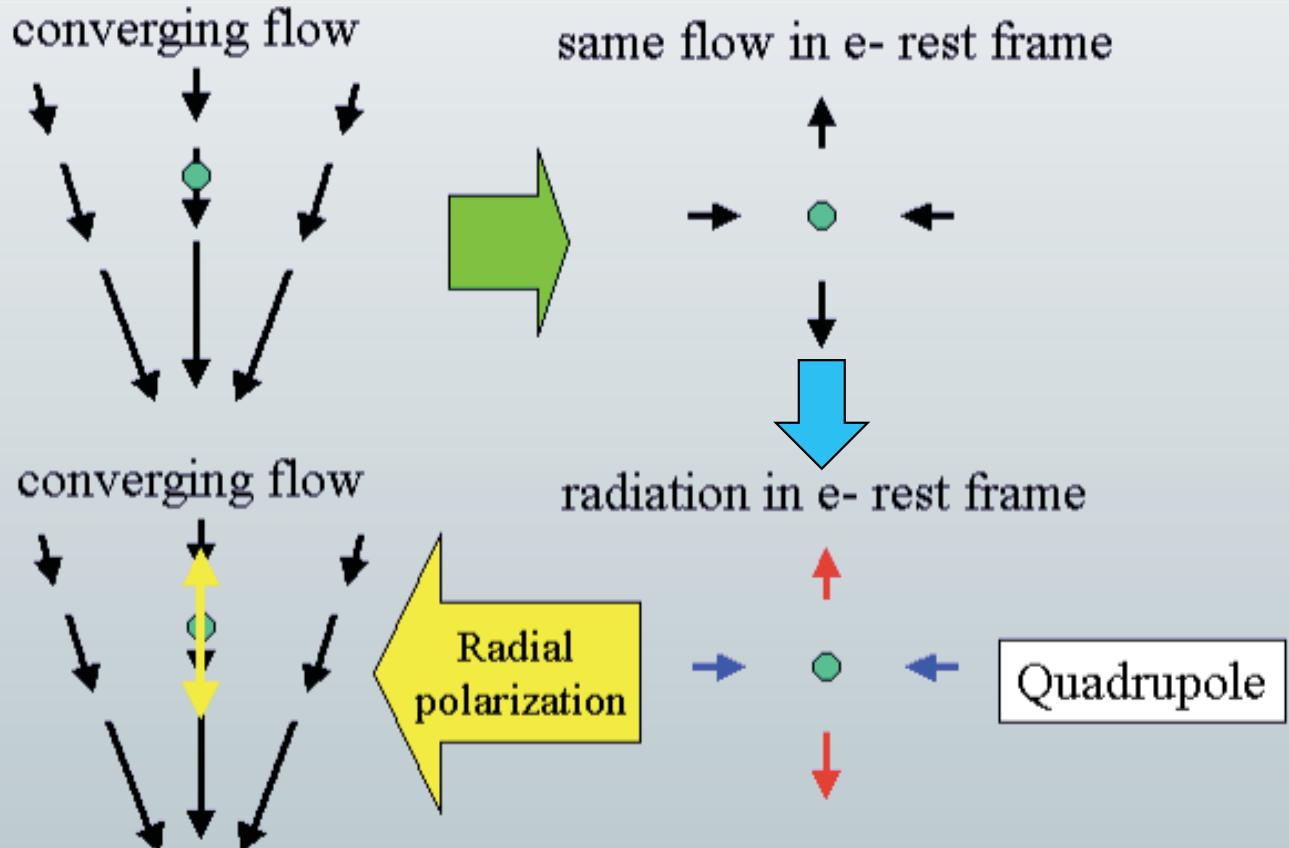


Mais pas de non gaussianité à petite échelle (fnl, gnl...)

# Thomson scatterings are polarised



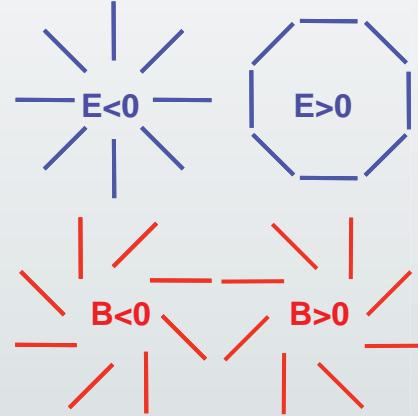
- Before recombination, successive scatterings destroy polarization and the radiation arrives at recombination unpolarized
- During recombination, Gradients in the velocity field can produce a quadrupole in the rest frame of the scattering electron



- A *diverging* flow leads to a *tangential* pattern of polarisation

# CMB angular power spectra

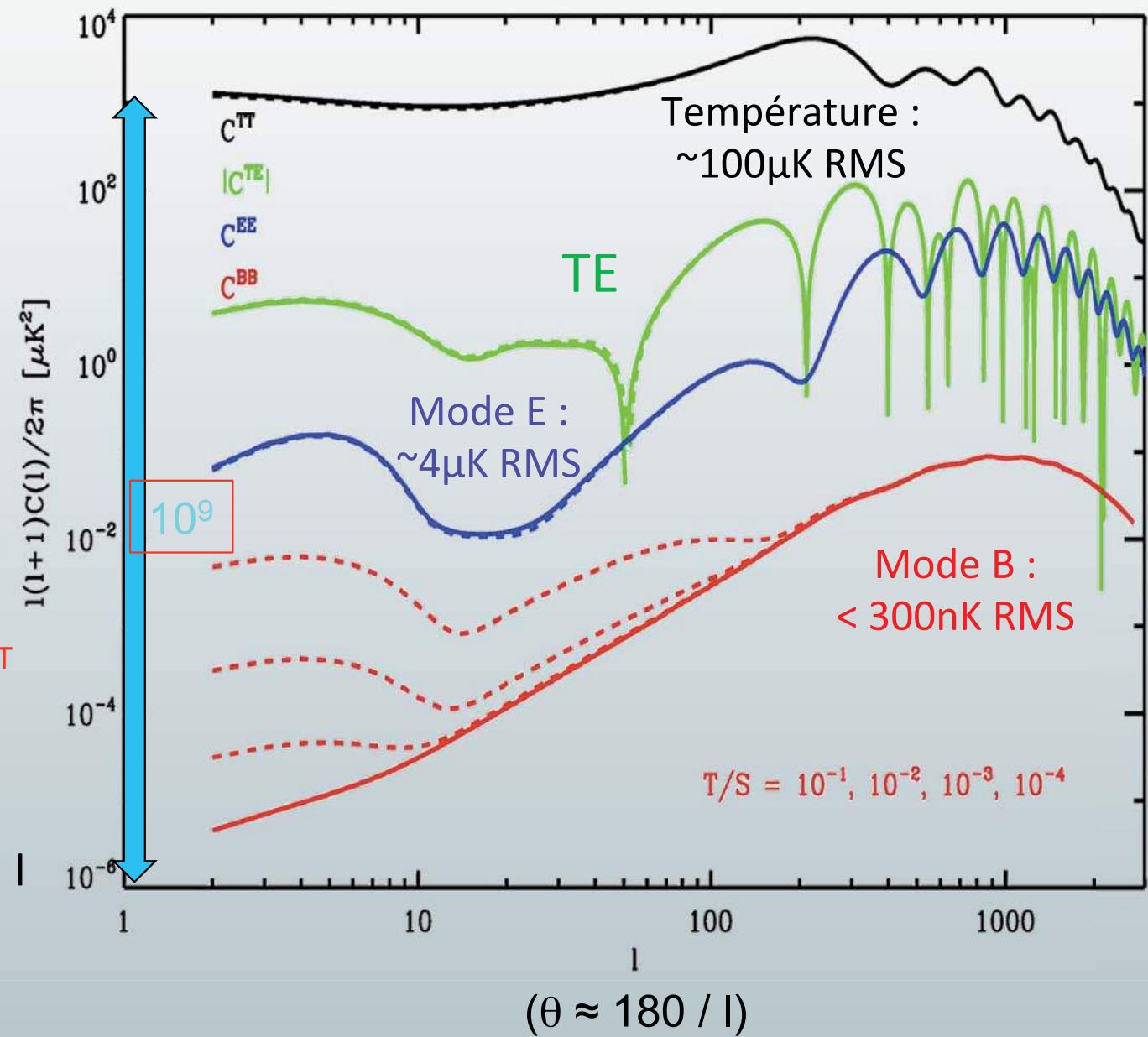
- 3 observables : T, E, B

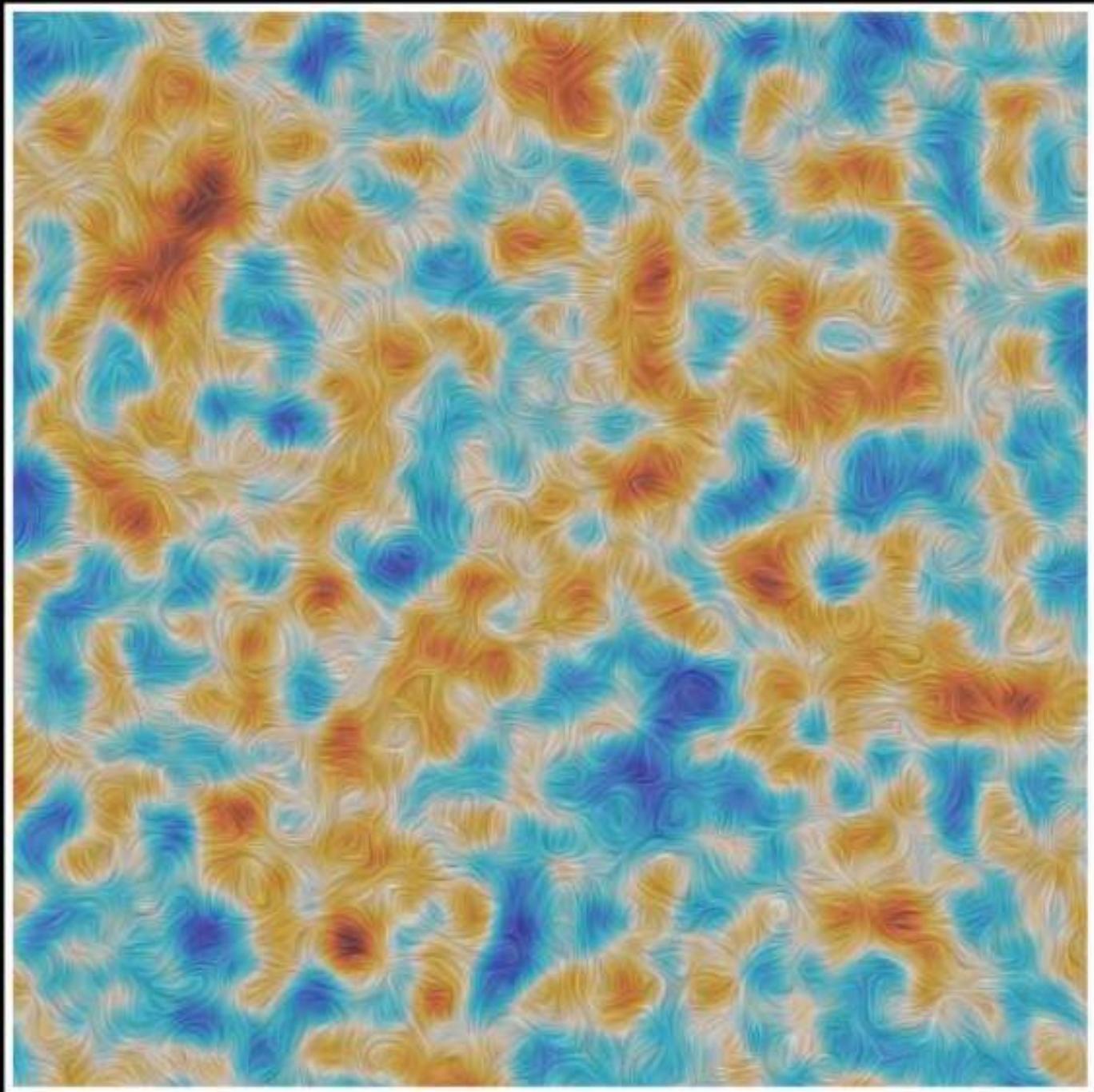


- B Modes:

- Not generated by scalar modes
- “Smoking gun” of tensorial perturbations
- At best 300 times weaker than T fluctuations
- case  $T/S = r = 0.1$  (cf. fig),
- $E_{inf} = 2 \times 10^{16}$  GeV (GUT).

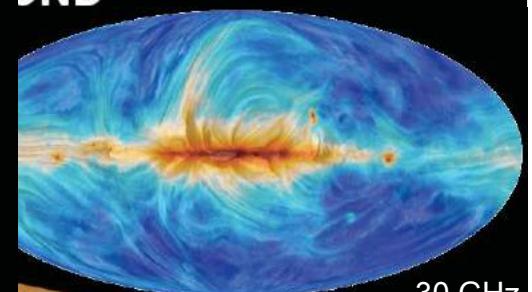
- B mode Spectrum peaks at  $l < 200$ , i.e.  $\theta > 1$  deg



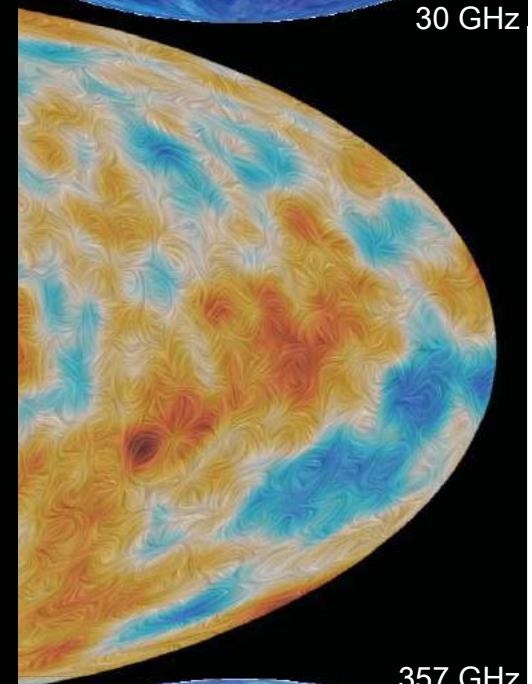


Filtered at 20 arcminutes

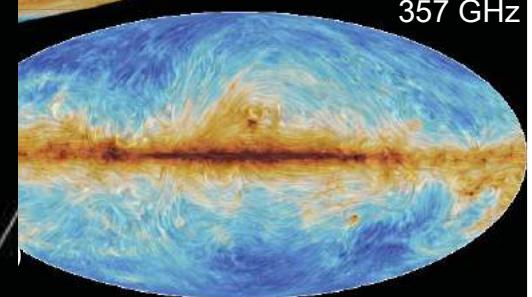
JND



30 GHz

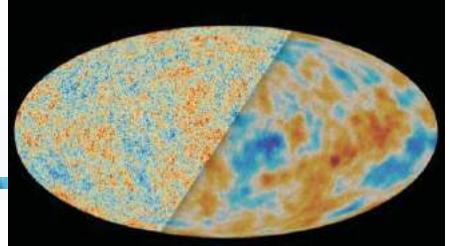


357 GHz

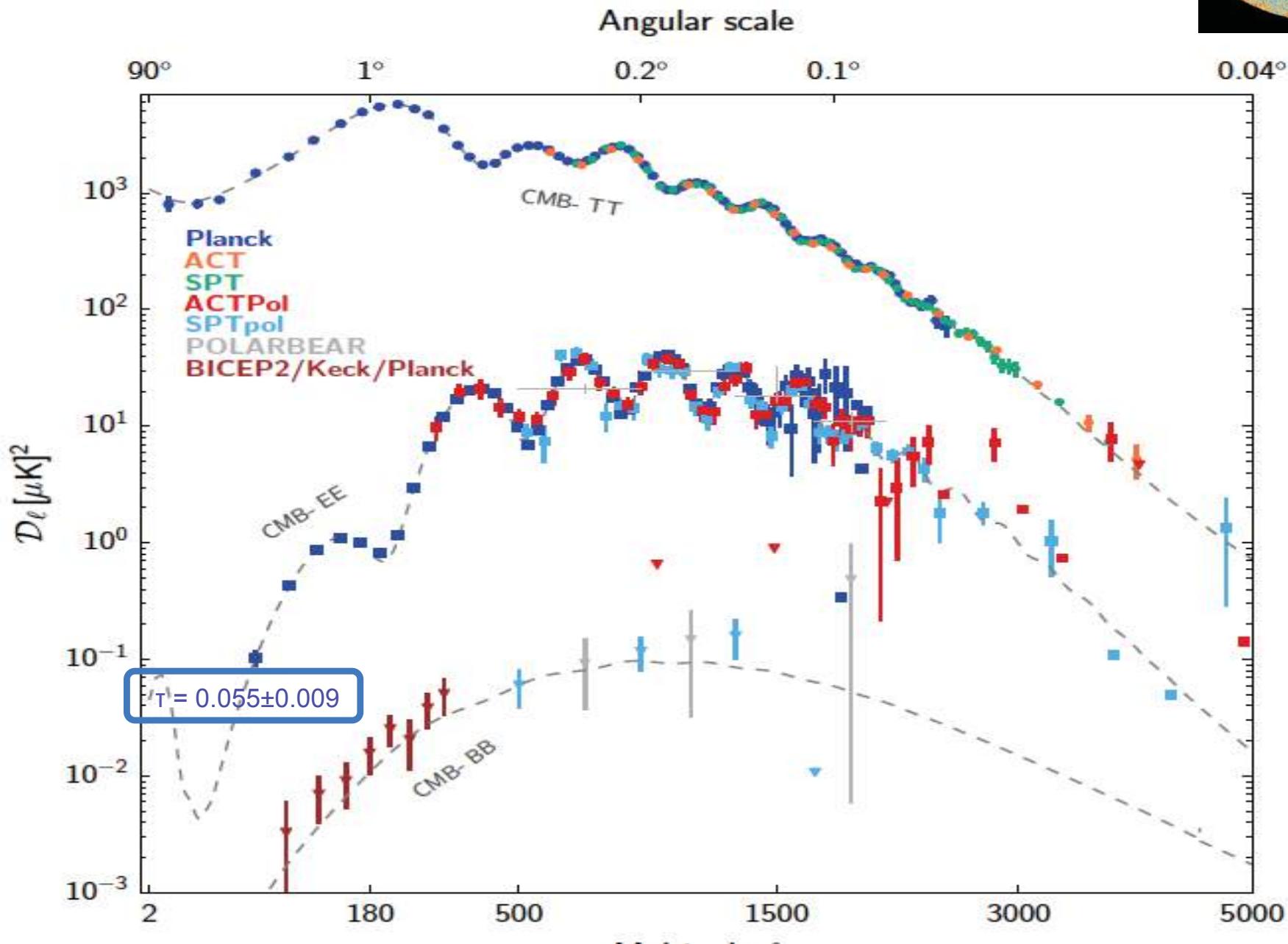


105

# TT, EE, BB – mid 2015 status



Only keeping points w. sufficiently small error bars, Fig. calabrese



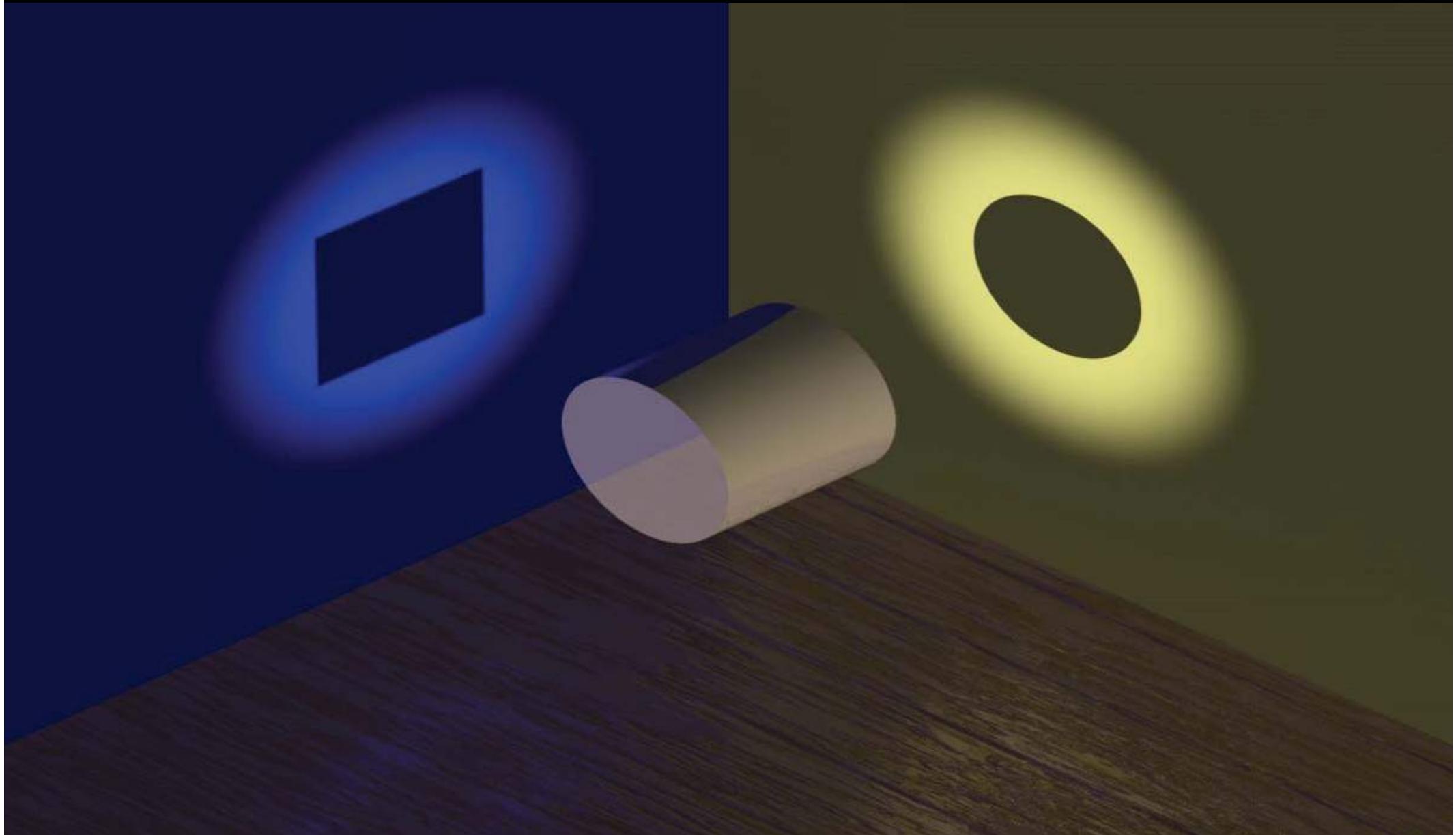
1 114 000 Modes measured with TT,  
60 000 with TE  
96 000 with EE  
... and 10's in BB  
+ weak constraints with TB and EB

# Base $\Lambda$ CDM model

Parameter	[1] <i>Planck</i> TT+lowP	[2] <i>Planck</i> TE+lowP
$\Omega_b h^2$ . . . . .	$0.02222 \pm 0.00023$	$0.02228 \pm 0.00025$
$\Omega_c h^2$ . . . . .	$0.1197 \pm 0.0022$	$0.1187 \pm 0.0021$
$100\theta_{\text{MC}}$ . . . . .	$1.04085 \pm 0.00047$	$1.04094 \pm 0.00051$
$\tau$ . . . . .	$0.078 \pm 0.019$	$0.053 \pm 0.019$
$\ln(10^{10} A_s)$ . . . . .	$3.089 \pm 0.036$	$3.031 \pm 0.041$
$n_s$ . . . . .	$0.9655 \pm 0.0062$	$0.965 \pm 0.012$
$H_0$ . . . . .	$67.31 \pm 0.96$	$67.73 \pm 0.92$
$\Omega_m$ . . . . .	$0.315 \pm 0.013$	$0.300 \pm 0.012$
$\sigma_8$ . . . . .	$0.829 \pm 0.014$	$0.802 \pm 0.018$
$10^9 A_s e^{-2\tau}$ . . . . .	$1.880 \pm 0.014$	$1.865 \pm 0.019$

TT & TE have similar uncertainties,  
but beware that they are still some low level systematics in the polarisation data

# Il aurait pu en être autrement!



And it further constrains potential deviations from the base tilted LCDM model/physics

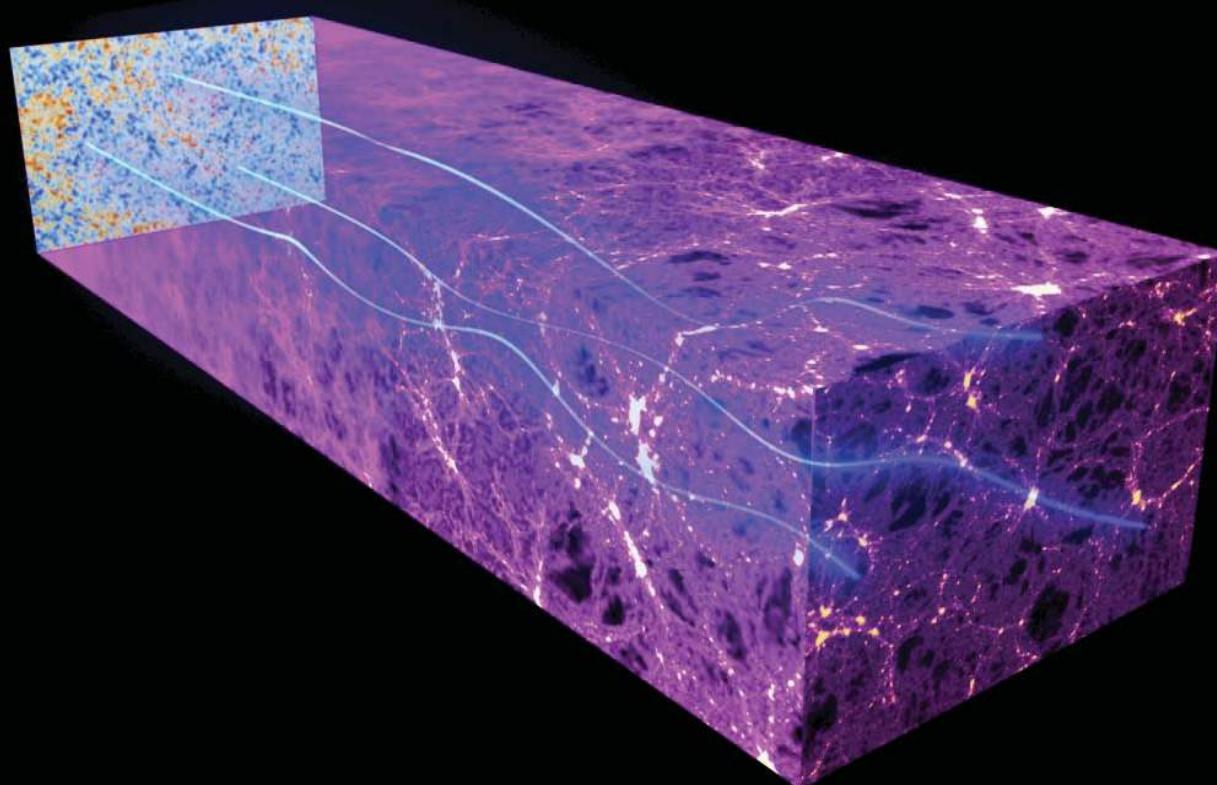
# Le modèle LCDM à 6 paramètres

- Un modèle élégant, incroyablement minimal,
  - à la simplicité trompeuse,  
*puisque'il repose sur nos deux théories fondamentales, la relativité générale et la mécanique quantique, et des hypothèses hardies, telle que*
    - *Les lois de la physique sont toujours et partout les mêmes,*
    - *L'univers à grande échelle est homogène et isotrope,*
    - *La relativité générale peut s'utiliser à des échelles beaucoup plus vastes qu'à celles où elle a été testée directement, une extrapolation spectaculaire par 15 ordres de grandeur...*
    - *Idem pour la mécanique quantique*
    - *Tout en utilisant la physique classique avec une précision horlogère*
- ... toute choses que nous testons aujourd'hui !*

# GRAVITATIONAL LENSING DISTORTS IMAGES

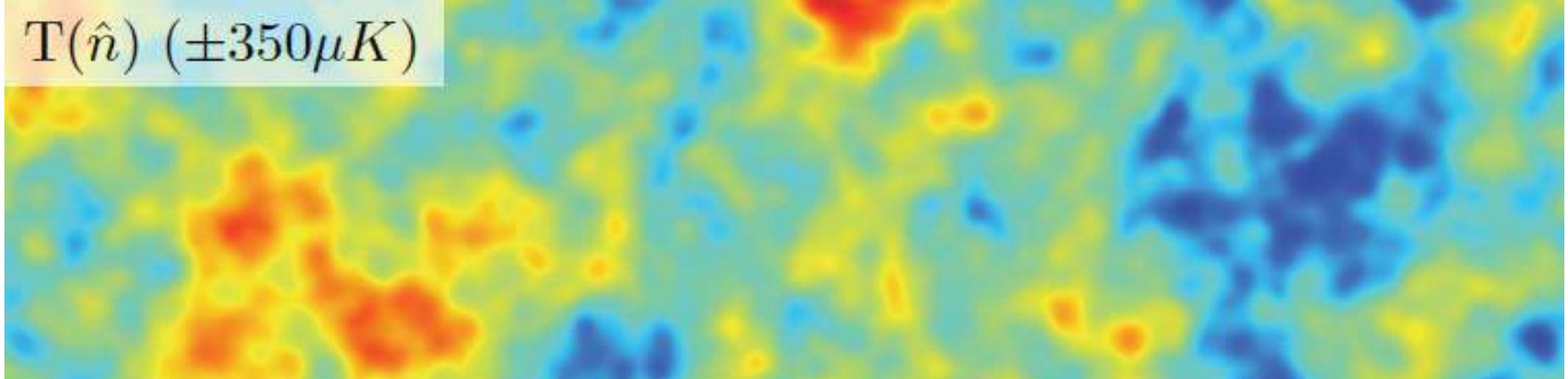


The gravitational effects of intervening matter bend the path of CMB light on its way from the early universe to the Planck telescope. This “gravitational lensing” distorts our image of the CMB (smoothing on the power spectrum, and correlations between scales)

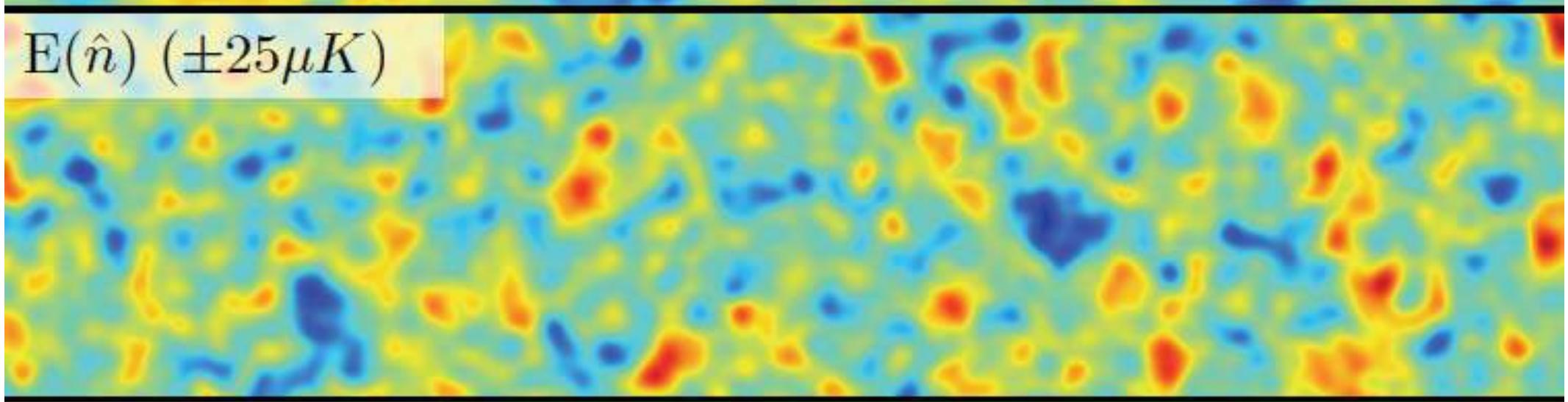


$$\begin{aligned}\hat{T}(\vec{\theta}) &= T(\vec{\theta} + \vec{\nabla}\phi) \approx T(\vec{\theta}) + \vec{\nabla}\phi \cdot \vec{\nabla}T(\vec{\theta}) + \dots \\ \bar{\phi} &= \Delta^{-1} \vec{\nabla} \cdot [C^{-1} T \vec{\nabla}(C^{-1} T)]\end{aligned}$$

$T(\hat{n})$  ( $\pm 350\mu K$ )



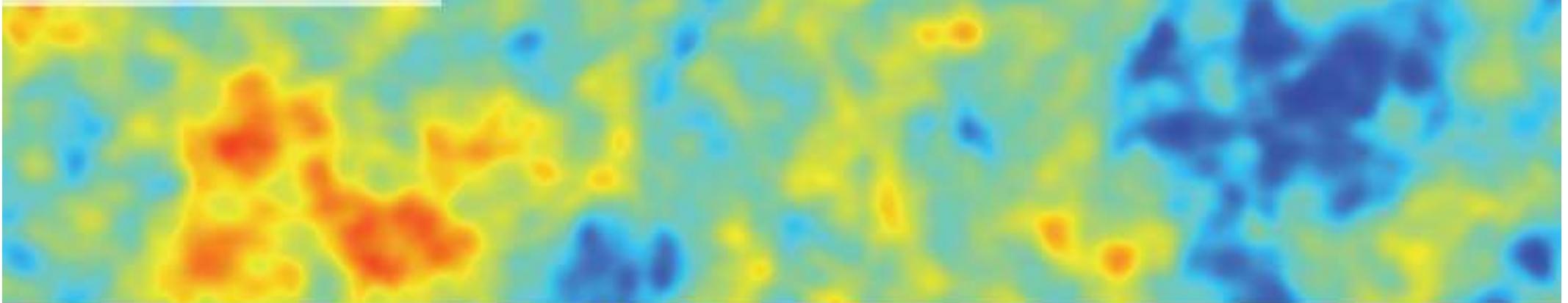
$E(\hat{n})$  ( $\pm 25\mu K$ )



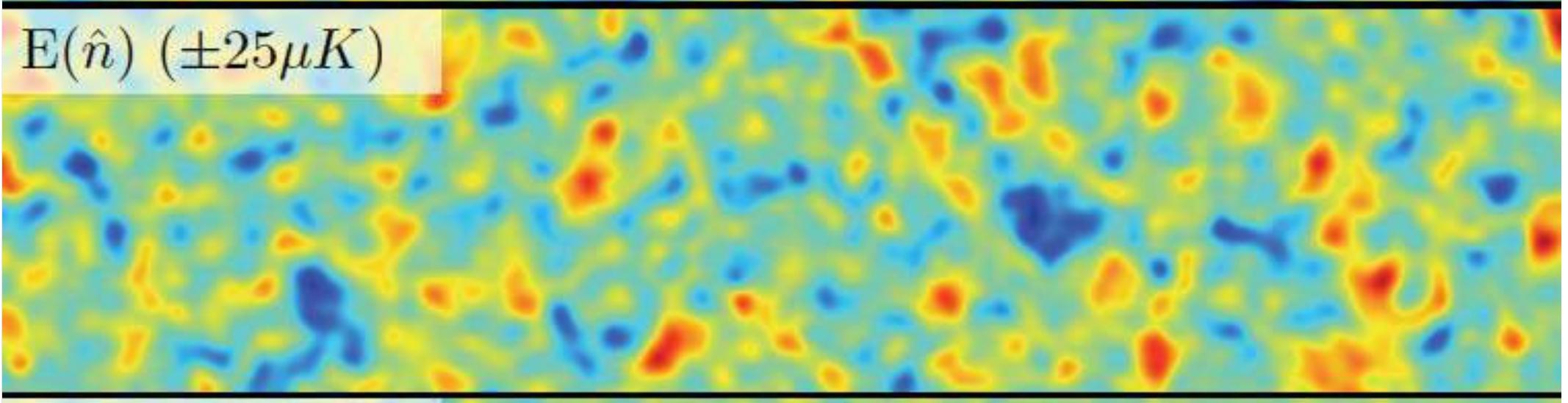
$B(\hat{n})$  ( $\pm 2.5\mu K$ )



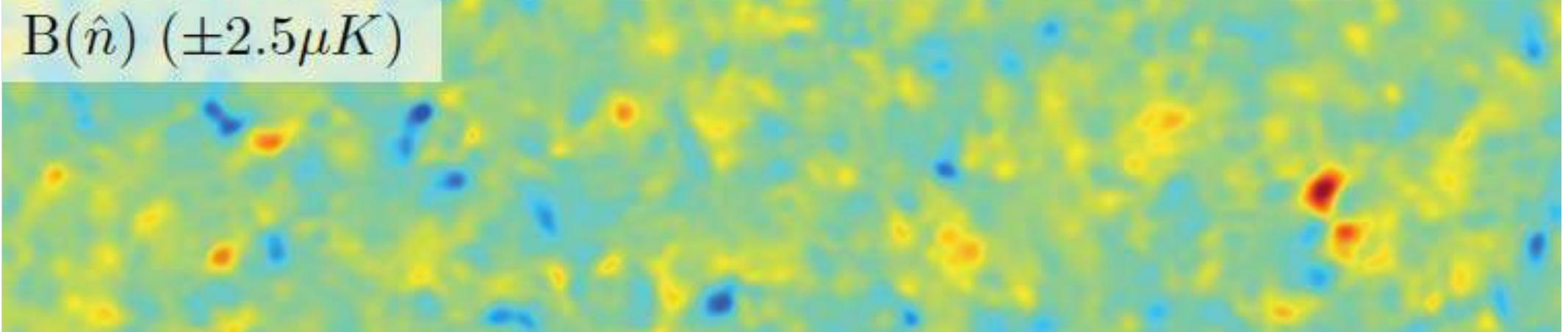
$T(\hat{n})$  ( $\pm 350\mu K$ )



$E(\hat{n})$  ( $\pm 25\mu K$ )



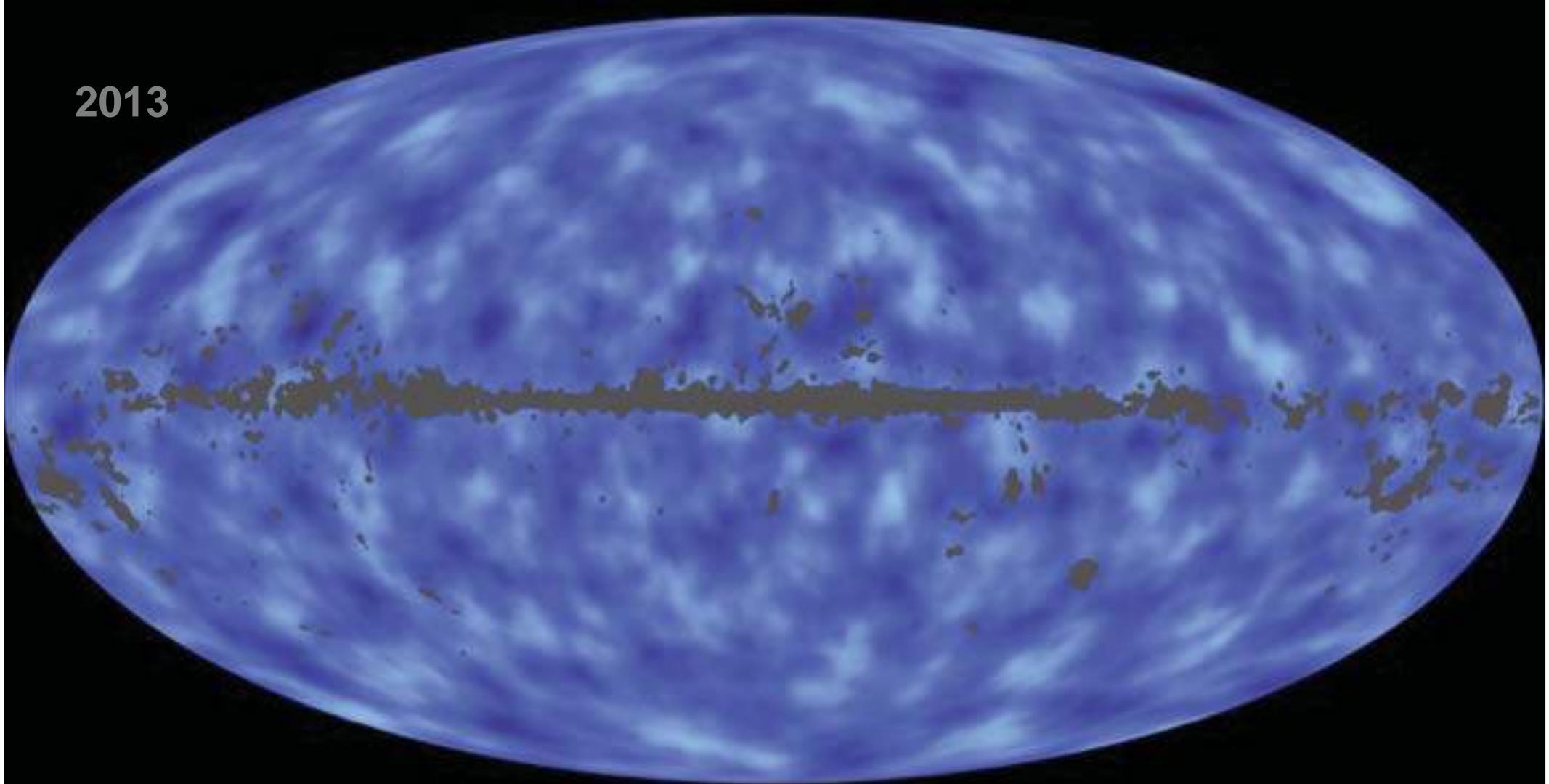
$B(\hat{n})$  ( $\pm 2.5\mu K$ )



# Projected mass map



2013

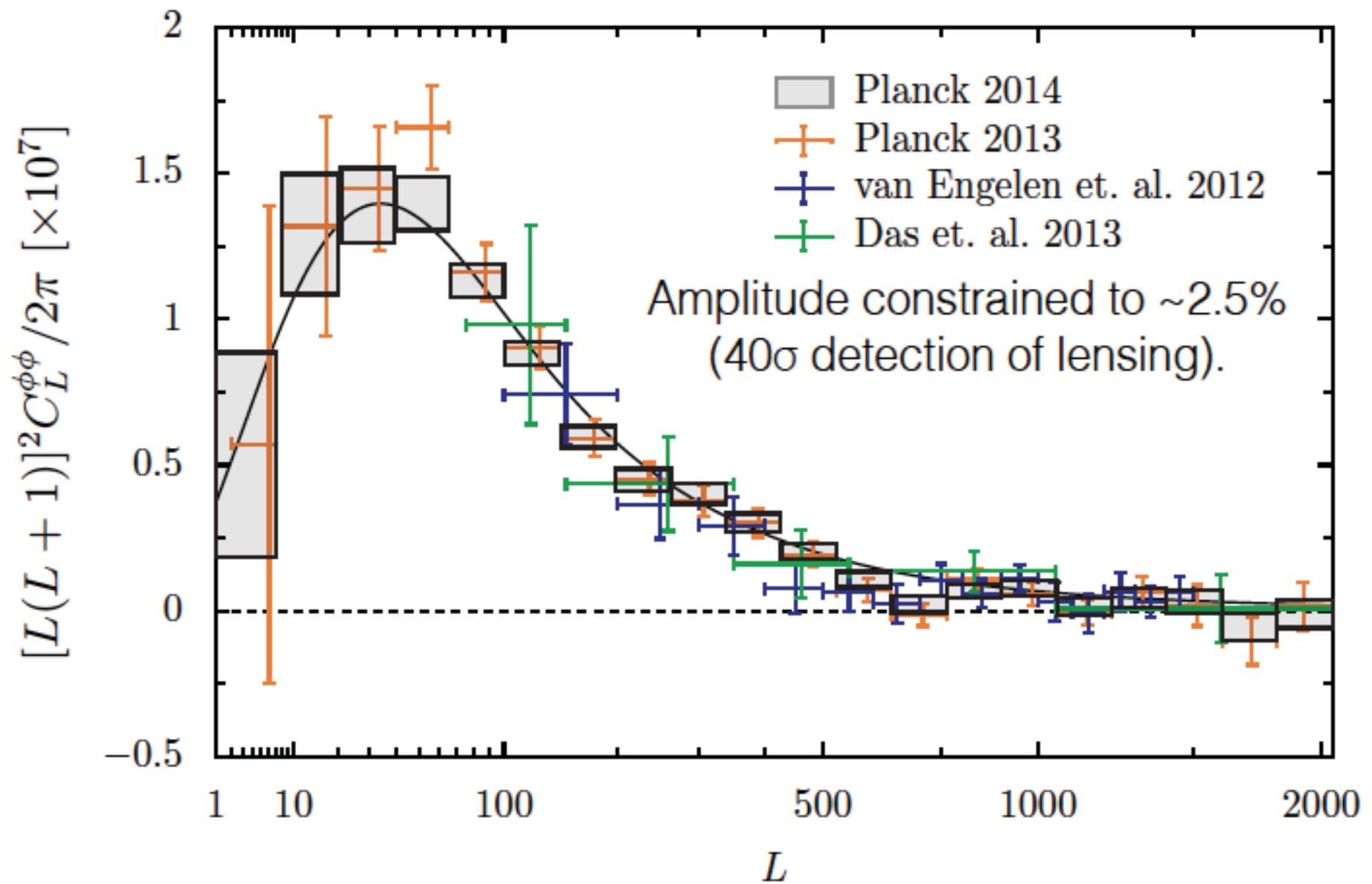


The (grey) masked area is where foregrounds are too strong to allow an accurate reconstruction

Page 140

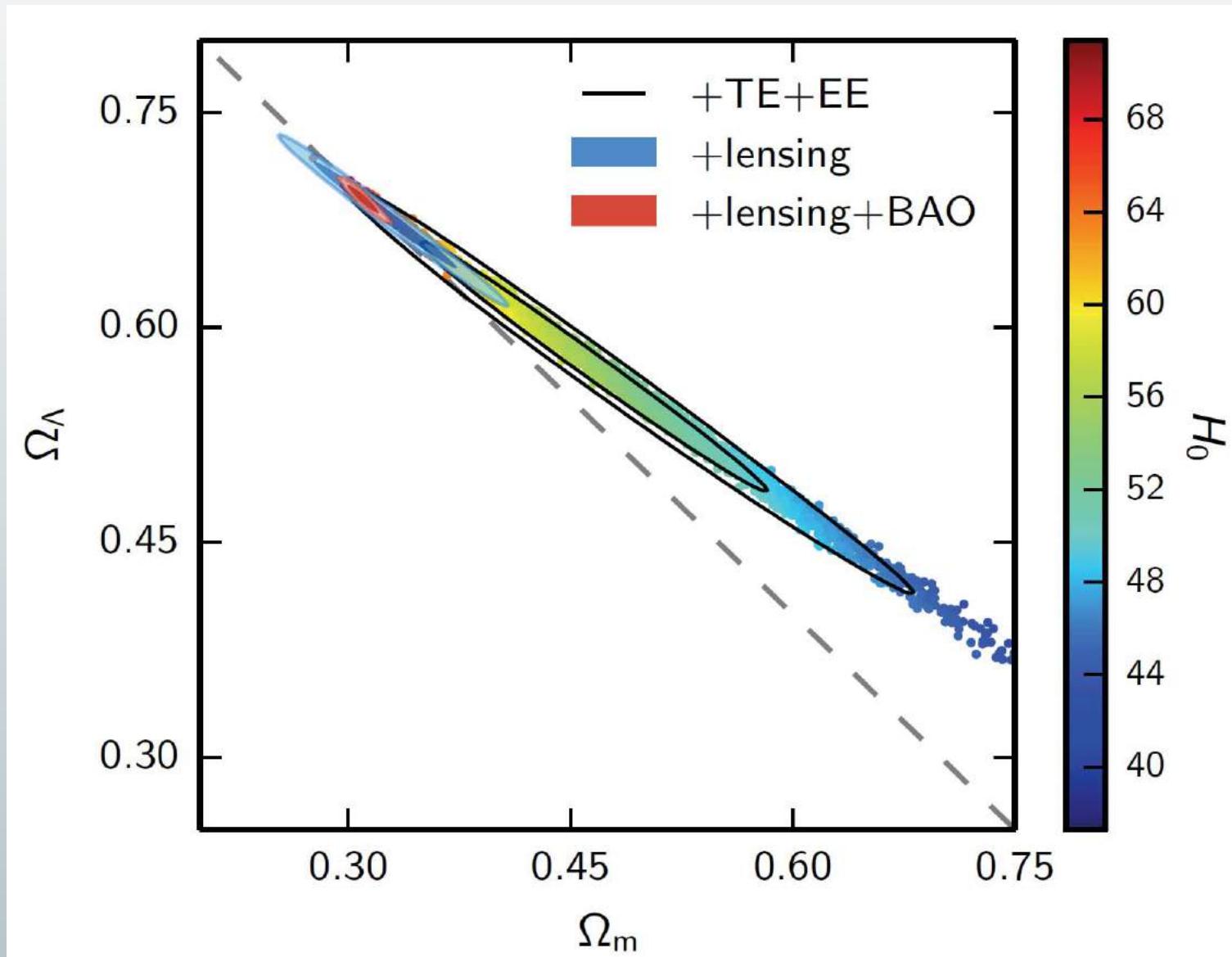
European Space Agency

# Lensing power spectrum

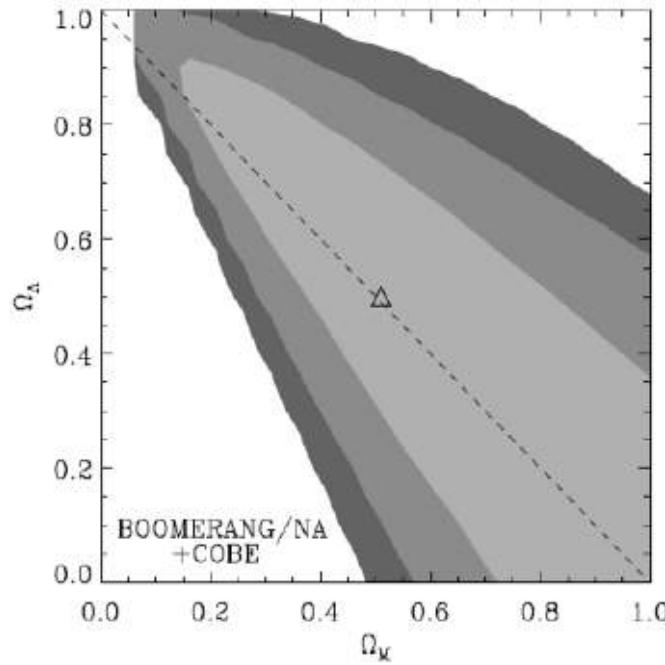


Planck for the first time measured the lensing power spectrum with higher accuracy than it is predicted by the base CDM model that fits the temperature data

# Spatial curvature constraint

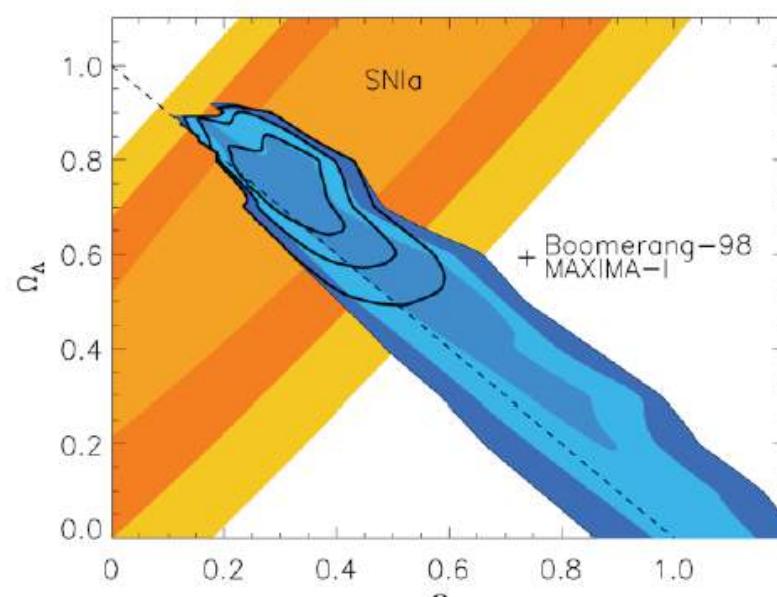


# Spatial curvature constraint



$$\Omega_K = -0.05^{+0.40}_{-0.40}$$

Melchiorri et al. 2000



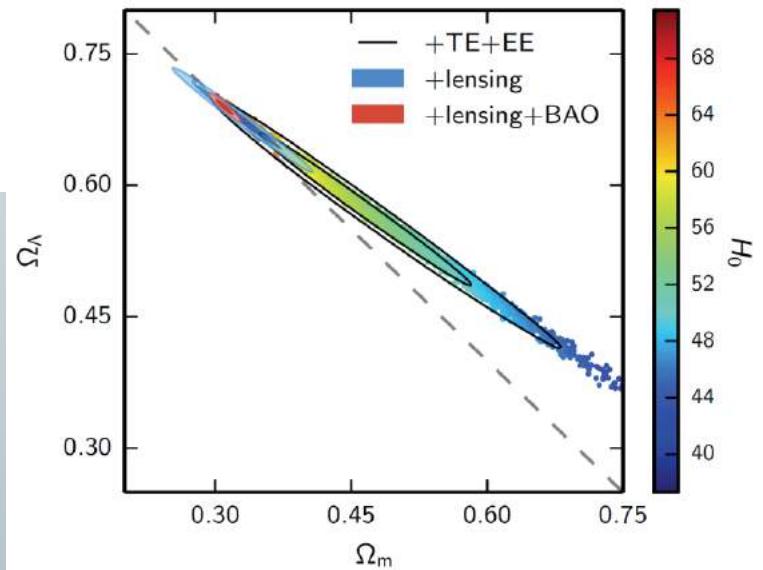
$$\Omega_K = -0.11^{+0.07}_{-0.07}$$

Jaffe et al. 2001

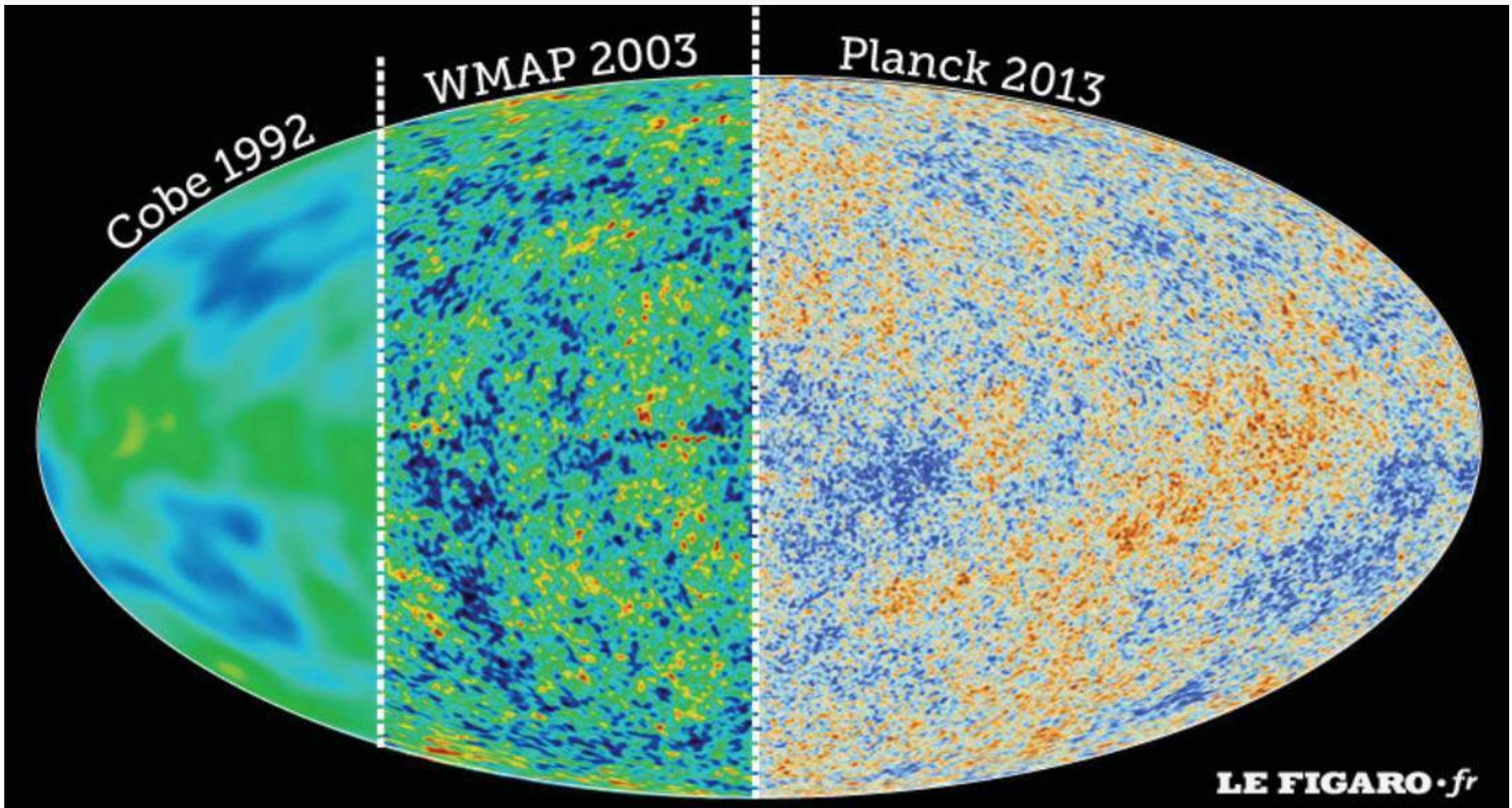
Planck 2015

$$\Omega_k = 0.000 \pm 0.005 \text{ (95% CL)}$$

Note the change of axes  
For Planck below



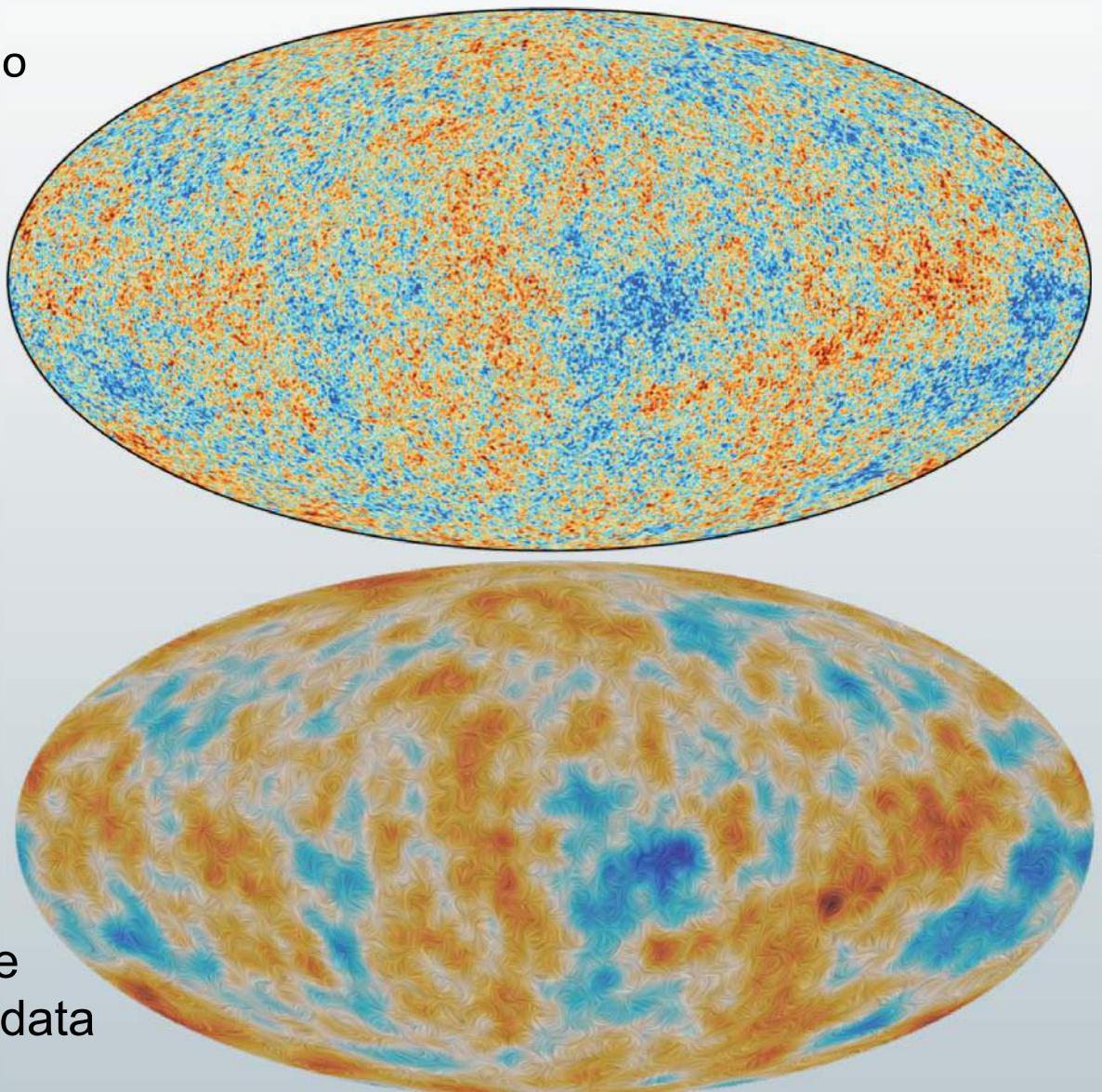
# Que de chemin parcouru...



# Conclusions

→ base  $\Lambda$ CDM continues to be a good fit to the Planck data, *including polarisation*.

Parameter	<i>Planck TT,TE,EE+lowP</i>
$\Omega_b h^2$ .....	$0.02225 \pm 0.00016$
$\Omega_c h^2$ .....	$0.1198 \pm 0.0015$
$100\theta_{\text{MC}}$ .....	$1.04077 \pm 0.00032$
$\tau$ .....	$0.079 \pm 0.017$
$\ln(10^{10} A_s)$ .....	$3.094 \pm 0.034$
$n_s$ .....	$0.9645 \pm 0.0049$
$H_0$ .....	$67.27 \pm 0.66$
$\Omega_m$ .....	$0.3156 \pm 0.0091$
$\sigma_8$ .....	$0.831 \pm 0.013$
$10^9 A_s e^{-2\tau}$ .....	$1.882 \pm 0.012$



→ powerful evidence in favour of simple inflationary models, that match Planck data to very high precision.

→ If there is new physics beyond base LCDM, then the corresponding observational signatures in the CMB are weak and difficult to detect.





HFI PLANCK

# Un regard vers l'origine de l'Univers

Physicsworld  
**TOP 10  
BREAKTHROUGH  
2013**

LA MISSION

LE SATELLITE

RÉSULTATS

LES PLANCIENS

NOTRE UNIVERS

OUTILS

Recherche...



## ACTUALITÉS

02 février 2015

### La réponse de BICEP2/KECK/Planck

En mars 2014, avec l'annonce de la détection d'un signal "prouvant" l'existence d'une phase d'inflation primordiale dans l'histoire de l'Univers, la collaboration BICEP2 avait fait la Une des médias du monde entier.

## LA RÉPONSE DE BICEP2/KECK/PLANCK

RÉSULTAT



## NOUVEAUTÉS ÉDUCATIVES



La séparation de composante, un vrai jeu d'enfant

Niveau : dès le CM1  
Durée : 5 à 15 min



Dessert cosmique

Niveau : dès le CM2

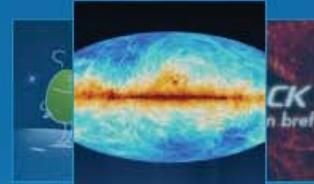
Toutes les malettes

## AGENDA

Aucune date

Voir l'agenda

## NOUVEAUX MÉDIAS



Exploration de notre galaxie avec Planck  
04/02/2015 - Cartes du ciel

Tous les médias



PLANCK EXPRESS



JEUX



ESPACE ÉDUCA



ESPACE PRESSE



JOURNAL DE BORD



MULTI MÉDIA



Nous suivre sur  
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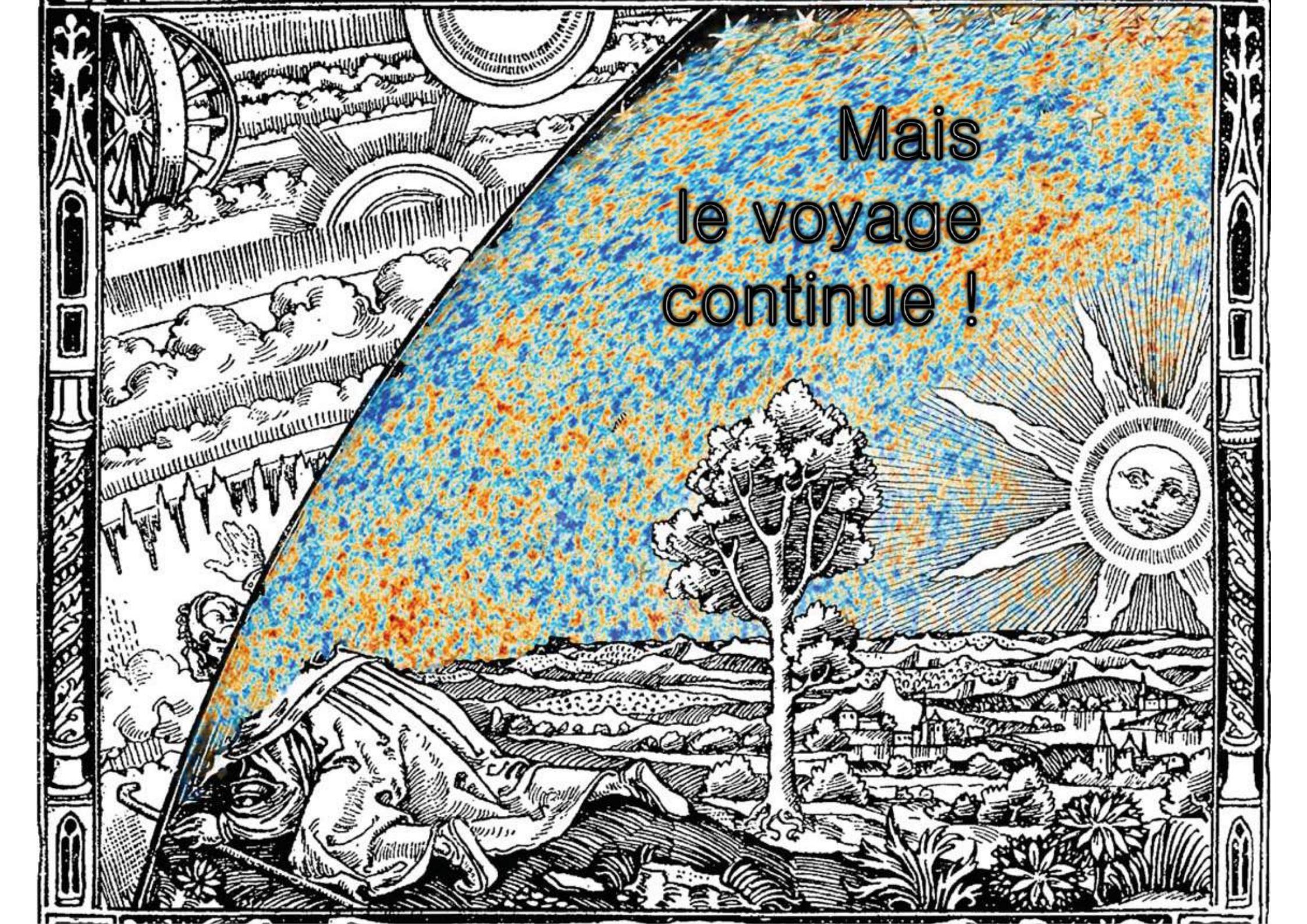


PLANCK EXPO

## PARTENAIRES

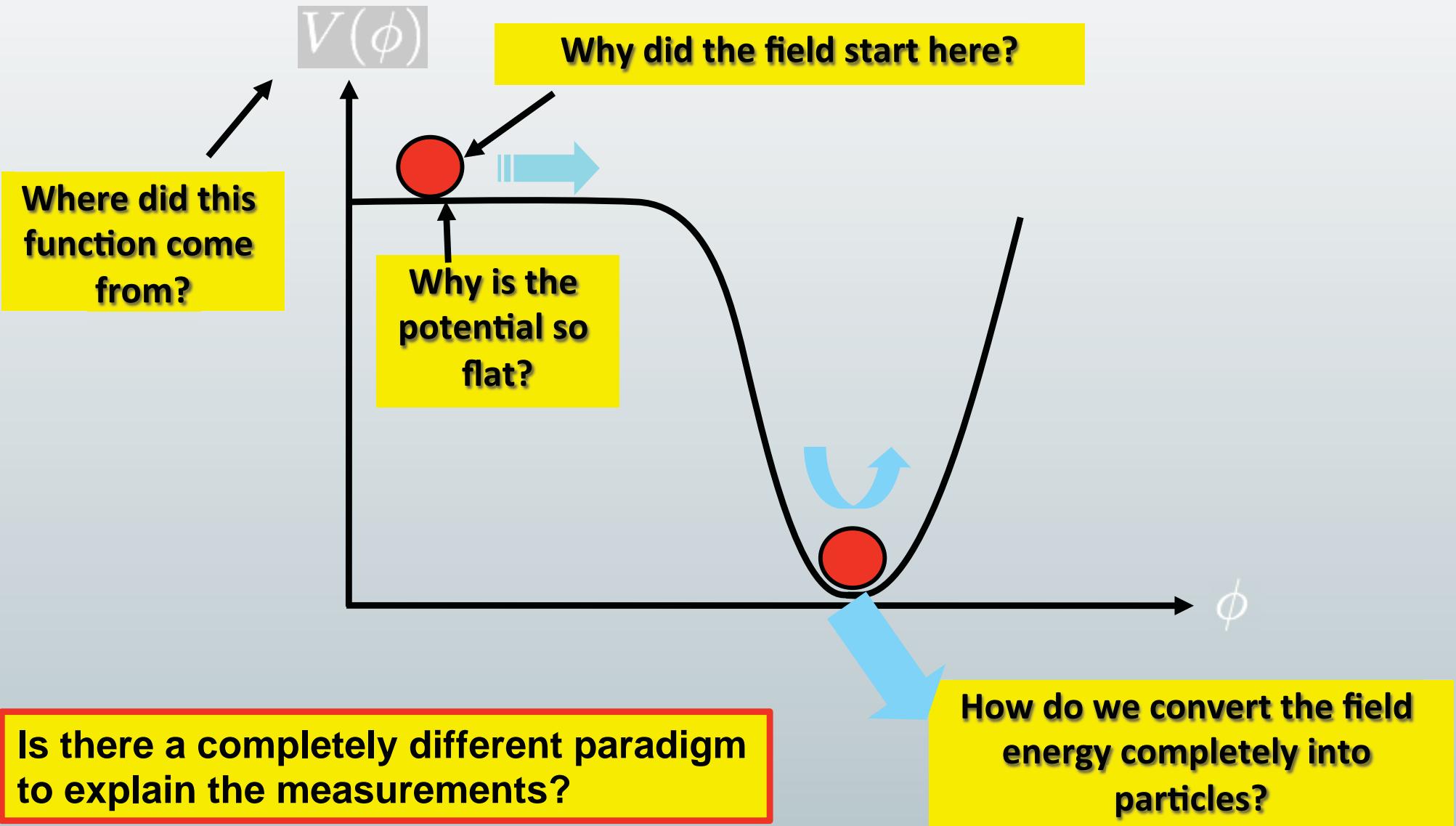


ThalesAlenia  
Space

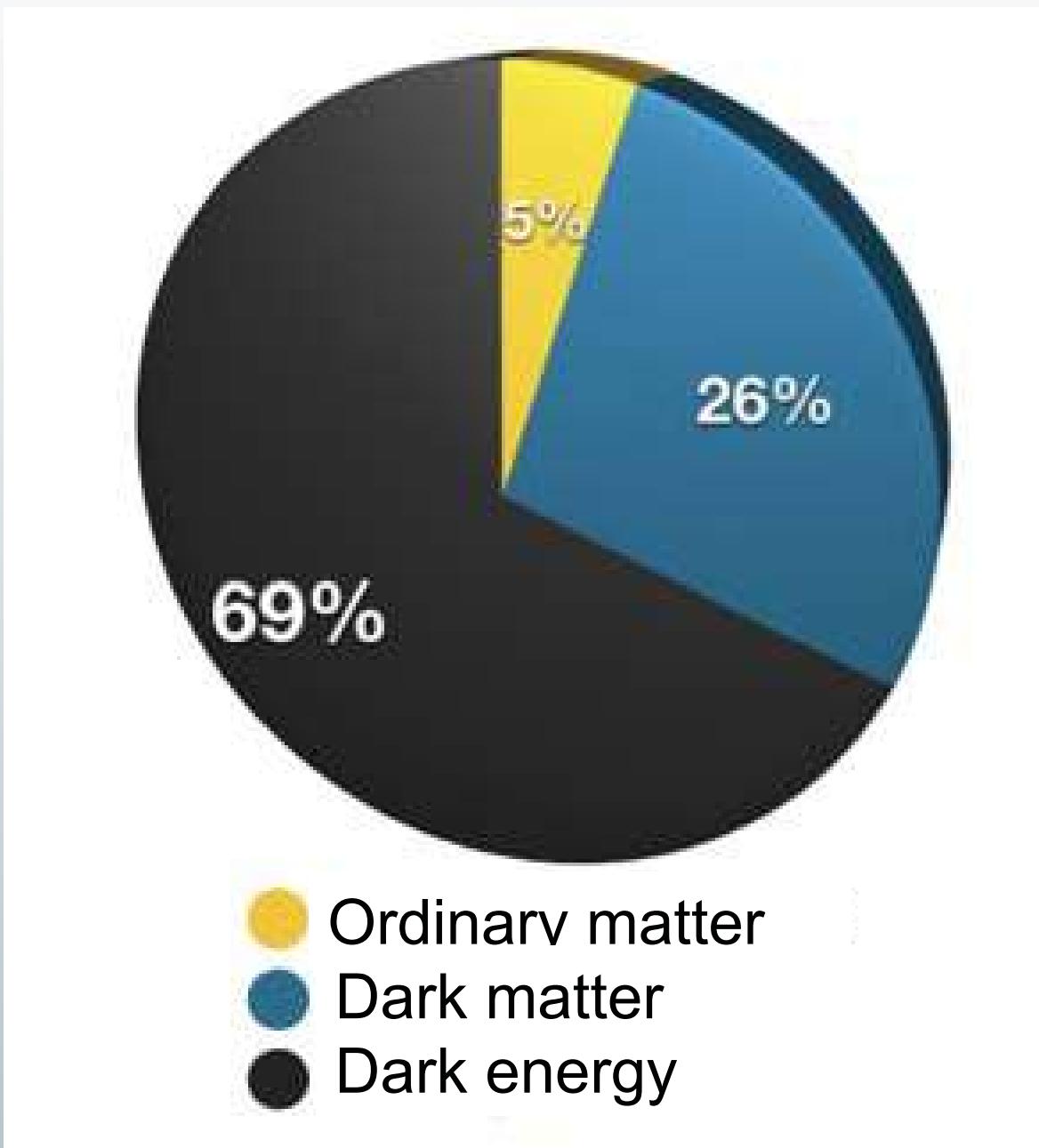


Mais  
le voyage  
continue !

# Indeed what is the physics of inflation?



# An extraordinary content today





**Thank you**